

Detailed spectroscopic analysis of the Trapezium cluster stars inside the Orion Nebula

Rotational velocities, stellar parameters and oxygen abundances ^{*}

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Abstract. We present the results of a spectroscopic analysis of the Trapezium cluster stars inside the Orion Nebula. The rotational velocities have been obtained using Fourier analysis method, finding agreement with values derived from the usual method, based on linewidth measurements. The rotational velocity derived for θ^1 Ori C through this method is consistent with the variability of some of its spectral features that have a period of 15.42 days. By means of the fit of H, He I and He II observed profiles with FASTWIND synthetic profiles, stellar parameters and wind characteristics have been derived. This methodology let us estimate the errors associated with these parameters. It is found that macroturbulence effects have to be included for a good fit to the He I-II lines in the spectrum of θ^1 Ori C.

By means of a very accurate study, oxygen abundances have been derived for the three B0.5V stars θ^1 Ori A, D and θ^2 Ori B. Final abundances are consistent with the nebular gas-phase results presented in Esteban et al. (2004) and are lower than those given by Cunha & Lambert (1994). Our results suggest a lower dust depletion factor of oxygen than previous estimations for the Orion nebula.

Key words. Stars: abundances – Stars: early-type – Stars: fundamental parameters – Stars: atmospheres – Stars: individual: θ^1 Ori C – Stars: rotation – ISM: abundances – ISM: H II regions – ISM: individual: Orion nebula

1. Introduction

New developments of massive star model atmosphere codes have led to interesting new possibilities for stellar spectroscopic studies. Improvements in computational methods as well as an increase of the efficiency of computers, have made possible the modeling of atmospheres of hot luminous stars, taking into account not only strong NLTE effects and hundreds thousands metallic lines producing the so-called *line-blanketing* (Hubeny & Lanz 1995), but also winds with expanding spherical geometries (Santolaya-Rey et al. 1997; Hillier & Miller 1998; Pauldrach et al. 2001; Puls et al. 2005).

The new improvements that have been included in these latest generation models call for a review of previous

results. For example, the papers by Herrero et al. (2002), Crowther et al. (2002), and Martins et al. (2002) showed that the SpT - T_{eff} calibrations used previously (Vacca et al. 1996) needed to be revised to lower effective temperatures for a given spectral type. Recent analyses by Repolust et al. (2004) and Martins et al. (2005) in the Milky Way, and by Massey et al. (2004, 2005) at SMC and LMC metallicities, reinforce this result and together implies a need to revisit the ionizing flux distribution that is used for the study of H II regions and starbursts.

Recent work by Trundle et al. (2002) and Urbaneja et al. (2005) have shown that abundance gradients in some spiral galaxies derived from stellar and nebular studies tend to be coherent but very dependent on the calibration used in the strong line nebular methods. However, until now, there have been no detailed systematic studies which compare results from nebular and stellar studies. This is the first of a series of papers

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^{*} The INT is operated on the island of La Palma by the RGO in the Spanish Observatorium of El Roque de los Muchachos of the Instituto de Astrofísica de Canarias.

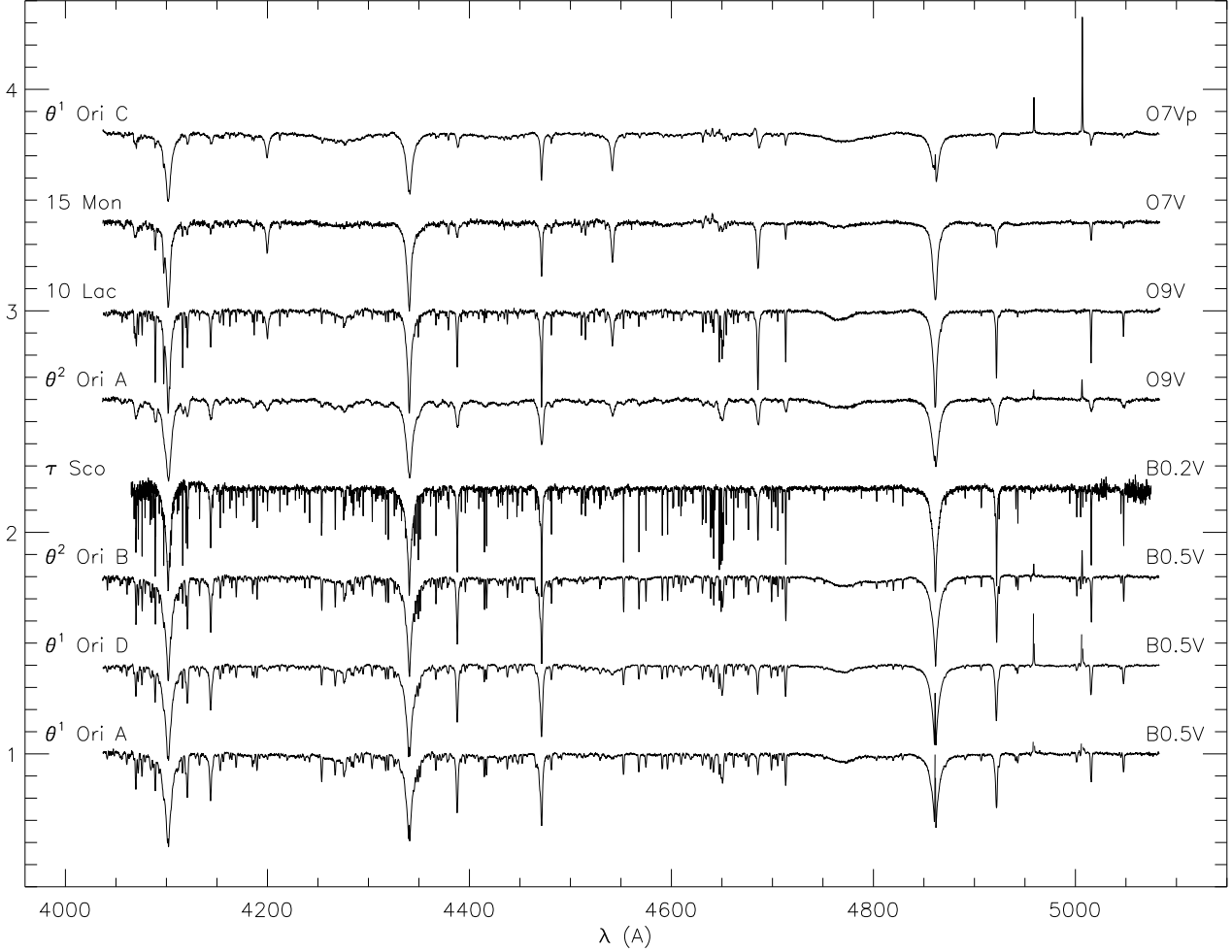


Fig. 1. Atlas of the INT+IDS spectra in the 4000-5000 Å region used for this study. τ Sco spectrum is from CASPEC. Several narrow nebular emission lines appear in the spectra of the Orion stars. In these stars, the cores of the hydrogen Balmer lines are contaminated by the nebular lines. Nebular H_β and $[O III] \lambda\lambda 4960, 5007$ lines have been arbitrarily diminished in the θ^1 Ori A and D plotted spectra for clarity reasons. Many metal lines can be distinguished together with the broad H and He lines in the O9 - B0.5 spectra. The high rotational velocity of θ^2 Ori A makes the metal lines of this O9V star shallower, broader and blended. The known inverted P-Cygni $He II \lambda 4686$ profile can be seen in the θ^1 Ori C spectrum.

aimed at this subject. We have selected some galactic H II regions for a detailed study of the interaction between massive stars and the surrounding ISM, looking for consistency of derived parameters (T_{eff} , luminosities and ionising flux distribution of the stellar population) as well as abundances of C, N, O, Si and Mg. For this first study we have selected the Orion nebula, a well studied and resolved Galactic H II region powered by a cluster of a few massive stars, the Trapezium cluster.

The Orion complex contains the massive on-going star forming region closest to Earth, at about only 450 pc. The Orion nebula, M42, is part of this complex. It is a well known H II region (e.g. O'Dell 2001; Ferland 2001) ionised by the Trapezium cluster stars (θ^1 Ori), a group of early type stars located in the core of the nebula. Together with θ^1 Ori C (HD 37022, O7V), the main ionising source, we find other B0.5V stars, perfect

targets for a stellar abundance study.

The most recent study of the chemical composition of the Orion nebula has been presented by Esteban et al. (2004) who reanalysed the emission line spectrum of the nebula to determine the physical conditions and abundances of the ionised gas-phase. Cunha & Lambert (1992, 1994) included some of the Trapezium cluster stars in a survey of B-type stars in the Orion OB1 association. They presented a spectroscopic analysis of these stars for determining C, N, O, Si and Fe stellar abundances.

This paper is focused on the spectroscopic analysis of the Trapezium cluster stars for deriving their stellar parameters and oxygen abundances. The stellar parameters obtained for θ^1 Ori C will be used in future papers as input for the modeling of the Orion Nebula with photoionization codes. The derived stellar abundances are compared to those obtained by Esteban et al. (2004)

through nebular studies.

Our paper is structured as follows: In Sect. 2 we present the observations. In Sect. 3 and 4 we obtain the $v \sin i$ and the stellar parameters of our targets. The study of θ^1 Ori C is presented in Sect. 5, and the oxygen abundance analyses in Sect. 6. A discussion of the results and the conclusions of this work are presented in Sect. 7.

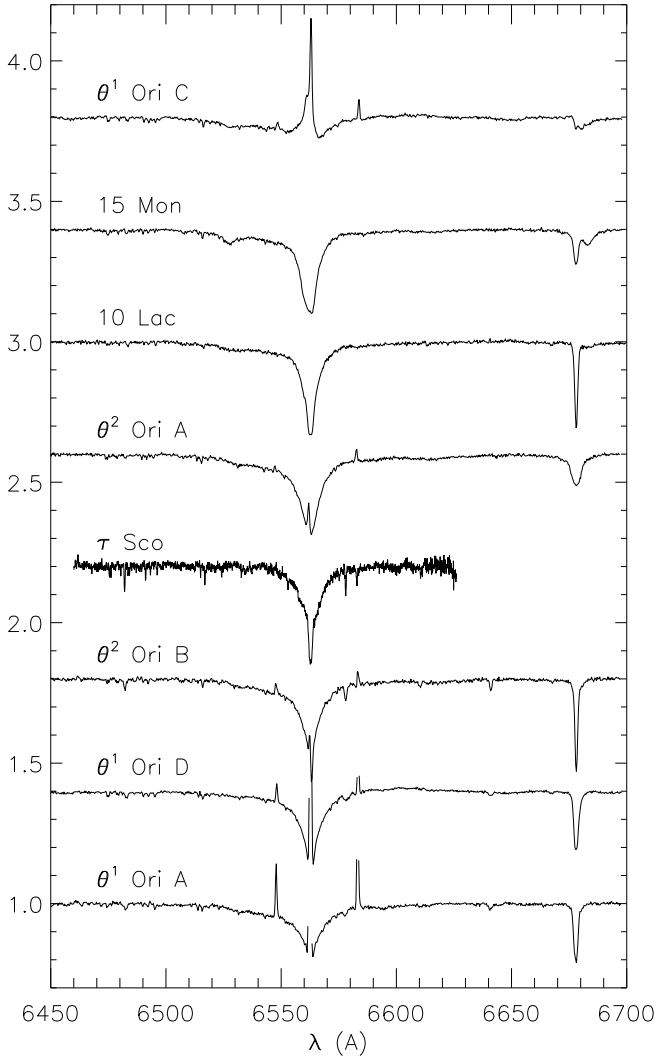


Fig. 2. Atlas of the INT+IDS spectra in the H_α region used for this study. τ Sco spectrum is from CASPEC. Nebular H_α and $[N II]$ emission lines have been diminished in the θ^1 Ori A and D plotted spectra for clarity reasons. Two emission components can be distinguished in the θ^1 Ori C H_α line: the narrow one is associated with the nebula, the other one is variable and it is associated with the star (see section 5). The blue wing of the H_α line is affected by some cosmetic in the INT observations.

2. The observations

The bulk of the observations used here were carried out with the Isaac Newton 2.5m. Telescope (INT) at the

Roque de los Muchachos observatory in La Palma on the 21th of December 2002. The Intermediate Dispersion Spectrograph (IDS) was used with the 235 mm camera and two different gratings. We have observed the spectral region between 4000 and 5050 Å using the H2400B grating, which resulted in an effective spectral resolution $R \sim 7500$ (equivalent to a 0.23 Å/pixel resolution and $\sim 2.6 \text{ pixel FWHM arc lines}$). For the H_α region the H1800V grating was used, resulting in a similar spectral resolution (0.3 Å/pixel , $R \sim 8000$). With the configurations used three exposures were needed to cover the whole range. A large number of flat fields and arcs for the reduction process were obtained.

The reduction and normalization of the spectra was made following standard techniques, with IRAF and our own software developed in IDL. The signal-to-noise ratio (SNR) of the reduced spectra depends on the spectral range, but is usually about 200-250 in the blue region and 250 in the H_α region (see Table 2).

We found some problem when rectifying the INT observations in the H_α region. In the blue wing of the H_α line we observed a feature which is independent of rectification. No known line should be present at this wavelength, so we argue that it must be cosmetic.

Special care has to be taken over the nebular contamination of the stellar spectra. The stars studied are located inside $H II$ regions, so the stellar spectra are contaminated by the nebular emission spectrum. It would be desirable to remove this nebular contribution, as it fills the cores of the Balmer H and He I lines. However this is not easy even though we have long slit observations; the nebular emission has spatial structure that complicates its subtraction (especially for the Balmer H lines, the most important nebular contribution). If this is not correctly done an over or under subtraction will appear. After trying different possibilities we concluded that the best solution for this problem is not to subtract the emission lines and to ignore these regions in the final spectrum. For the Balmer H and He I lines this is satisfactory, as emission lines are narrower than absorption lines. Nebular contributions could be more difficult to separate for metal lines, however the contamination of the stellar metal lines used for the abundance analysis due to nebular lines is negligible.

The INT observations consist of the brightest three stars in the Trapezium cluster (θ^1 Ori A, C, D) together with the two nearby stars θ^2 Ori A and B. Two standard stars were included in this set, 10 Lac and 15 Mon (O9V and O7V respectively). The other standard star, τ Sco, was kindly provided by Dr. Gehren. This is a slow rotating B0.2V star perfect for a preliminary abundance analysis study. The spectrum was obtained with CASPEC, attached to the ESO 3.6m telescope. The SNR of this spectrum is ~ 200 -300 in the blue region and ~ 150 in the H_α region.

HD	Name	SpT	m_v	A_v	M_v
Orion stars					
HD 37020	θ^1 Ori A	B0.5V	6.73	1.89	-3.4
HD 37022	θ^1 Ori C	O7Vp	5.12	1.74	-4.9
HD 37023	θ^1 Ori D	B0.5V	6.71	1.79	-3.3
HD 37041	θ^2 Ori A	O9V	5.07	1.12	-4.3
HD 37042	θ^2 Ori B	B0.5V	6.41	0.73	-2.6
Reference stars					
HD 47839	15 Mon	O7V	—	—	-4.8
HD 214680	10 Lac	O9V	—	—	-4.4
HD 149438	τ Sco	B0.2V	—	—	-3.3

Table 1. Identification, spectral type and photometric visual data of the studied objects. A_v and m_v values for Orion stars from Preibisch et al. (1999). M_v values for these stars have been calculated considering a distance $d \sim 450 \pm 50$ pc to the Orion nebula. Data for HD 214680 and HD 47839 are from Herrero et al. (1992). Photometric data for HD 149438 are from Humphreys (1978). Uncertainties in m_v , A_v and M_v are 0.01, 0.03 and 0.3 respectively.

HD	4000 - 5000	4600 - 5100	H_α
HD 37020	170-280	210	220
HD 37022	170-330	280	290
HD 37023	220-450	450	320
HD 37041	200-275	230	210
HD 37042	160-220	250	230
HD 214680	140-190	200	160
HD 47839	89-120	160	260

Table 2. *SNR* achieved for the different spectra for the three ranges observed with the INT+IDS.

For the study of the spectral variability of θ^1 Ori C, we have used FEROS spectra (some of them downloaded from the FEROS database and other kindly provided by Dr. Stahl). These observations were carried out with FEROS at the ESO 1.52m telescope in La Silla. The instrument is designed for high-dispersion spectroscopy with $R \sim 48000$ in the spectral range 3700 - 9200 Å. The achieved *SNR* is 300 at about 4500 - 5000 Å. Different phase observations have been used for the variability study (see Table 7). This study is presented in section 5.

3. Determination of rotational velocities

The analysis of stellar spectra makes use of a number of free parameters like the micro and macroturbulent velocities and the projected rotational velocity, $v \sin i$. The last one has acquired particular importance in recent times because of the mixing that rotation may induce in the interior of massive stars (e.g. Maeder & Meynet, 2000; Villamariz et al., 2002). However, some methods to determine the rotational velocities do not distinguish between rotation and other surface broadening mechanisms, like macroturbulence.

Conventionally, $v \sin i$ values are based on linewidth measurements of individual features, mainly metal lines apparently free of blends. As the principal broadening mechanism of these lines is stellar rotation, with sufficient resolution it is possible to determine $v \sin i$ from the FWHM of the line. Usually metal lines are used, however in cases of high rotational velocities or high temperatures metal lines appear blended or are very weak. Therefore, in these cases, if the $v \sin i$ is high, the whole He spectrum is used; if $v \sin i$ is not extremely high, only He I lines are used, as these lines are less affected by pressure broadening than He II lines. However, the $v \sin i$ derived must be tested with some metal lines (if available), as we are not completely sure of rotation broadening dominating over pressure broadening.

The Fourier method of determining $v \sin i$ is based on the fact that in the Fourier space convolutions transform into products and of the rotation, macroturbulence, natural and instrumental profiles (turbulence and instrumental assumed gaussian), only the rotation function has zeroes in its Fourier transform. These zeroes will appear in the total transform function, and Carroll (1933) showed that the position of the zeroes are related to the $v \sin i$. Actually the frequency of the first zero (σ_1) is related to the rotational velocity through:

$$\frac{\lambda}{c} v \sin i \sigma_1 = 0.660 \quad (1)$$

The microturbulence also introduces zeros in the Fourier transform at high frequencies (Gray, 1973). This has to be taken into account for very low $v \sin i$ ($\leq 20 \text{ km s}^{-1}$).

The main problems in the application of the Fourier method for early type stars are related to the quality of the observed spectra (i.e. spectral resolution and *SNR*). The lowest $v \sin i$ limit that can be determined is given by the spectral resolution ($\Delta\lambda$ in Å/pixel), as the sampling of the computational Fourier transform cannot be extended beyond $0.5/\Delta\lambda$. The noise in the observed spectra transforms as white noise that obscures the first zero.

The advantages of the Fourier method are that rotational broadening can be separated from other broadening mechanisms, and therefore metal lines as well as He lines can be used for the $v \sin i$ determination (even for low values of $v \sin i$). This is very useful for fast rotating stars and spectral types earlier than O9, showing blended or very weak metal lines.

We have tested the Fourier method with theoretical and observational cases, and it works well for massive hot stars. A paper with these results as well as determinations for a number of O star rotational velocities is in preparation (Simón-Díaz & Herrero, 2005), figure 3 shows a typical example. For a recent application to A-stars see Royer et al. (2002).

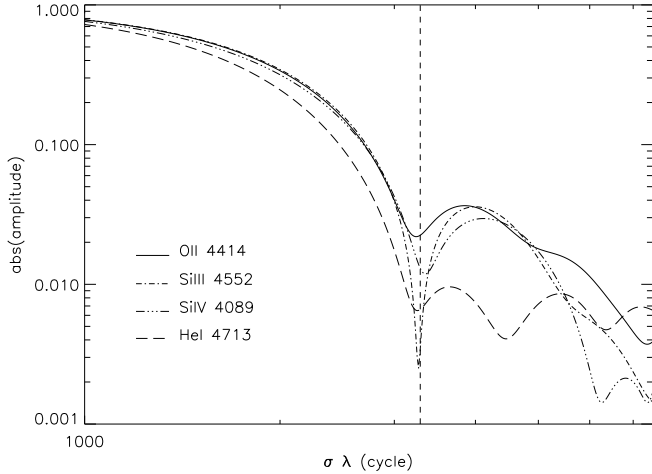


Fig. 3. Four different synthetic lines generated with FASTWIND have been convolved with a $v \sin i$ of 60 km s^{-1} and degraded to a SNR of 200. The different lines show the Fourier transform of the lines on a $\sigma\lambda$ baseline. The first zero is very close to the theoretical value for all the lines (even for the case of the He I line).

HD	$v \sin i \text{ (km s}^{-1}\text{)}$		
	This work		Other authors
	Fourier	FWHM	FWHM
HD 37020	55.0 ± 0.6	57 ± 3	56^a
HD 37022	24 ± 3	40 ± 11	53^b
HD 37023	49.0 ± 0.9	51 ± 3	51^a
HD 37041	131 ± 4	140 ± 11	110^b
HD 37042	34.0 ± 0.5	36 ± 3	10^c
HD 214680	30.0 ± 0.8	37 ± 4	35^b
HD 47839	67 ± 4	66 ± 6	67^b
HD 149438	< 13	16 ± 3	$5^d, 19^e$

Table 3. Projected rotational velocity derived from the Fourier and FWHM analyses. References from the literature are: ^a Simón-Díaz et al. 2003, ^b Howarth et al. 1997, ^c McNamara & Larsson 1962, ^d Schönberner et al. 1988, ^e Killian et al. 1991

The $v \sin i$ of our sample of stars has been determined through the Fourier and FWHM methods. Results are presented in Table 3, together with some values found in the literature, for comparison. The resolution in the IDS spectra is 0.23 Å/pixel , so the lowest $v \sin i$ that could be detected through Fourier method is $\sim 20 \text{ km s}^{-1}$. For θ^1 Ori C FEROS spectra the resolution is $\sim 0.03 \text{ Å/pixel}$, so $v \sin i \geq 2 \text{ km s}^{-1}$ for detection. For the τ Sco spectrum, the resolution is 0.1 Å/pixel , so the lowest detectable $v \sin i$ is 8 km s^{-1} .

All the papers referenced in Table 3 use the FWHM method applied to the optical spectra of the stars except those by Howarth et al. (1997), who use a cross correlation technique for IUE spectra, and Schönberner et al. (1988) who compare the observed spectrum of τ Sco with

synthetic profiles.

Comments on the individual stars' $v \sin i$ determination, as well as the comparison between the values derived through Fourier and FWHM methods are presented in Section 4.1. Agreement between both methodologies is very good, however there are some interesting cases (see the study of θ^1 Ori C in Section 5).

4. Stellar parameters

The analyses have been performed using the latest version of FASTWIND (an acronym for Fast Analysis of Stellar atmospheres with WINDs), a code which was originally described by Santolaya-Rey et al. (1997). See Puls et al. (2005) for the newest description of the code along with a discussion of comparisons with previous models and other spherical mass-losing codes. The latest version uses a more complete *line-blanketing* and a temperature correction method based on the energy balance of electrons. The technique used for the derivation of the stellar parameters is already standard and has been described elsewhere (Herrero et al. 2002; Repolust et al. 2004). We only give here the main points. The analyses are based on visual fitting of hydrogen Balmer lines and He I and He II lines. Through the He I/He II ionization equilibrium, the effective temperature can be estimated; the wings of the Balmer lines are useful for the determination of the gravity and can give us some information about the stellar wind.

The code also needs other parameters, such as microturbulence, He abundance and wind properties (mass loss, terminal velocity and β parameter). Actually wind properties are related through the Q parameter ($Q = \dot{M} / (v_\infty R)^{1.5}$).

Once the observed lines are fitted with the modeled ones, effective temperature, gravity, He abundance, microturbulence and $\log Q$ are determined. The low density in the winds of the studied objects makes the spectrum insensitive to changes in Q , so that we can only determine upper values in most cases. Microturbulence has no effect on the H/He spectrum of early type stars of large gravity, as it has been shown by Villamariz & Herrero (2002). Therefore only effective temperatures, gravities and He abundances are determined for this step of the analysis. Of course, microturbulence is important for the derivation of metallic abundances and will be determined in the corresponding section. The code also provides the emergent flux distribution and then mass, radius and luminosity can be calculated if M_v is known (see Herrero et al. 1992).

Errors in T_{eff} and $\log g$ can be established generating a microgrid around central values; visual comparisons between modelled lines and observations allow us to determine the range of possible values for these parameters. For further comments on the effects of varying the various physical parameters used in the analyses, their mutual interplay and their error bars see viz. Repolust

HD	Name	T_{eff} ± 1000 K	$\log g$ ± 0.1 dex	R (R_{\odot})	M (M_{\odot})	$\log(L/L_{\odot})$	$\log Q$
HD 37020	θ^1 Ori A	30000	4.0	6.3 ± 0.9	14 ± 5	4.45 ± 0.13	< -13.5
HD 37022	θ^1 Ori C	39000	4.1	9.9 ± 1.4	45 ± 16	5.31 ± 0.13	—
HD 37023	θ^1 Ori D	32000	4.2	5.6 ± 0.8	18 ± 6	4.47 ± 0.13	< -13.5
HD 37041	θ^2 Ori A	35000	4.1	8.2 ± 1.1	39 ± 14	4.93 ± 0.13	< -13.5
HD 37042	θ^2 Ori B	29000	4.1	4.5 ± 0.6	9 ± 3	4.11 ± 0.13	< -13.5
HD 214680	10 Lac	36000	3.9	8.3 ± 1.1	20 ± 7	5.02 ± 0.13	< -13.5
HD 47839	15 Mon	40000	4.1	9.3 ± 1.3	40 ± 14	5.30 ± 0.13	-13.0
HD 149438	τ Sco	32000	4.0	5.6 ± 0.8	11 ± 4	4.47 ± 0.13	-13.0

Table 4. Stellar parameters derived from FASTWIND analysis. Only an upper limit for $\log Q$ can be derived for these stars. The microturbulences considered for the HHe analysis in each star are shown in the corresponding fitting plots. A normal value for the He abundance has been considered for all the stars ($\epsilon = 0.09$).

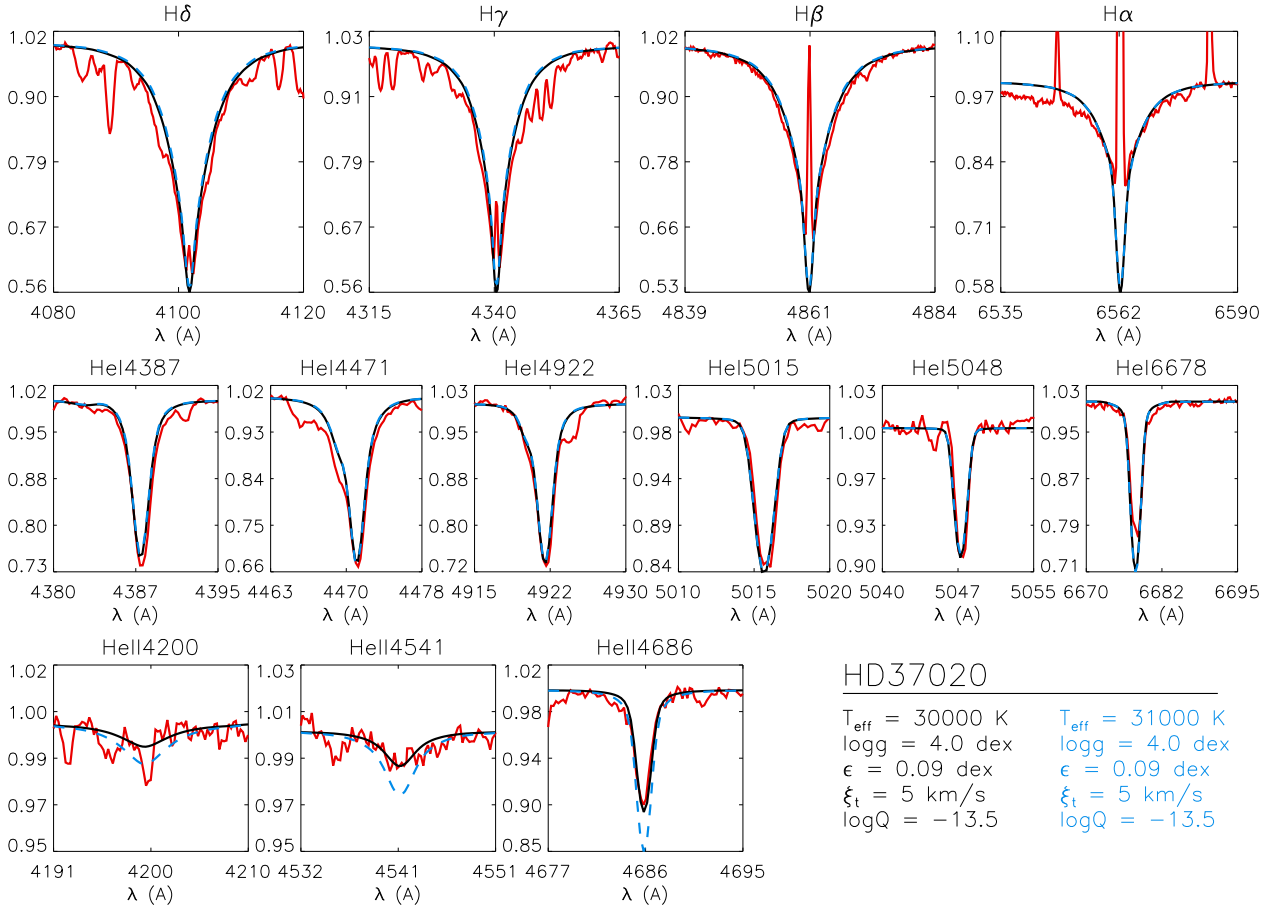


Fig. 4. Analysis of HD 37020 (θ^1 Ori A, B0.5V). Two models have been plotted for comparison with the observed spectrum. The adopted model has been represented with a solid line, while another model varying 1000 K in T_{eff} has been plotted with a dashed line. Note the core of the stellar Balmer lines contaminated with emission from the nebula. Only the wings of the Balmer lines are used for the fitting. The narrow line that appears in the core of the He II 4200 line is a N III absorption line; the red wing of the He I 4922 line is contaminated by an O II absorption line. The blue wing of the H α line is affected by a cosmetic feature and is not used for the fitting

et al. (2004), Herrero et al. (2002), Villamariz & Herrero (2002), and Villamariz et al. (2002). Errors in R , M and L are calculated considering the propagations of the uncertainties in T_{eff} , $\log g$ and M_{v} .

The fits of the synthetic FASTWIND H and He I-II

profiles to the observed ones are shown in Figures 4 to 10. The derived parameters for our sample of stars are shown in Table 4, corresponding to the best fits.

Some comments on the individual analyses and the comparison between spectroscopic and evolutionary results are given in Sects. 4.1 and 4.3.

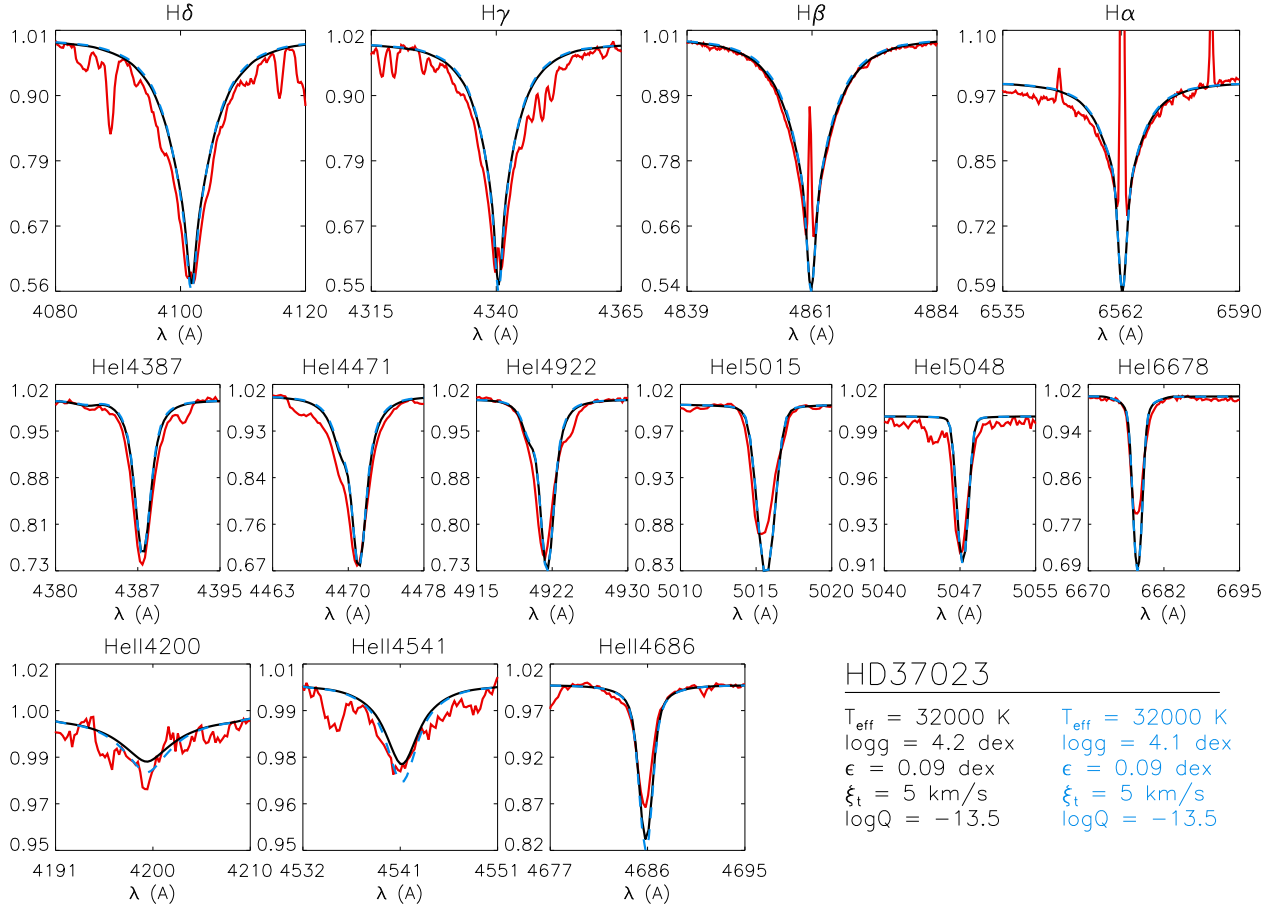


Fig. 5. As Fig. 4 for HD 37023 (θ^1 Ori D, B0.5V). A variation of 0.1 dex in $\log g$ has been considered in this case.

4.1. Comments on the Trapezium cluster stars

HD 37020: This is an eclipsing binary discovered by Lohsen (1975). The period is $P = 65.43233$ d, the magnitude range $m_v = 6.7 - 7.6$ and the eclipse lasts ~ 20 h. The primary is a B0.5V star which has many isolated metal lines. We have not found lines of the secondary star in our spectrum. This fact, together with the characteristics of the binary system described before, suggest that the secondary is a cooler, smaller and less luminous star, and that the change in the visual magnitude is due to the eclipse of the primary by its companion. Therefore the optical spectrum is mainly dominated by the B0.5V star.

The Fourier method has been applied to some O II, Si III-IV, N II and He I lines, deriving a $v \sin i = 55 \pm 0.6$ km s $^{-1}$. Metal and He I lines are in agreement. The $v \sin i$ derived through the linewidth measurement method is consistent with this value (see Table 3).

Figure 4 shows the good fit of the FASTWIND profiles to the observed spectra for the parameters given in Table 4 (except for the forbidden component of He I $\lambda 4471$ which is not well reproduced throughout our analyses. Note also that the apparent bad fit of

He II $\lambda 4200$ is due to the blend with the N III line at the same wavelength). We are very close to the applicability limit of the He I-II ionization equilibrium for deriving the T_{eff} as He I $\lambda 4200$ and He I $\lambda 4541$ lines are very faint; however these lines are still sufficiently sensitive to changes in T_{eff} and $\log g$ for deriving the stellar parameters accurately. The stellar parameters obtained here are in very good agreement with those obtained by Cunha & Lambert (1992) using the Strömgren index c_0 and the wings of H_γ with used Kurucz's (1979) LTE model atmospheres (this is not the case for the other two stars in common with these authors). A comparison of values is given in Table 5.

HD 37022: This is the main ionising source of the Orion nebula. See section 5 for a detailed study of this star.

HD 37023: This is the only star in the four Trapezium cluster stars (θ^1 Ori) without a binary companion (Preibisch et al. 1999). Robberto et al. (2004) find indications that this star is surrounded by a photoevaporated circumstellar disk.

The Fourier method gives a $v \sin i = 49.0 \pm 0.9$ km s $^{-1}$ for this star. Again, there is agreement with the linewidth measurement method.

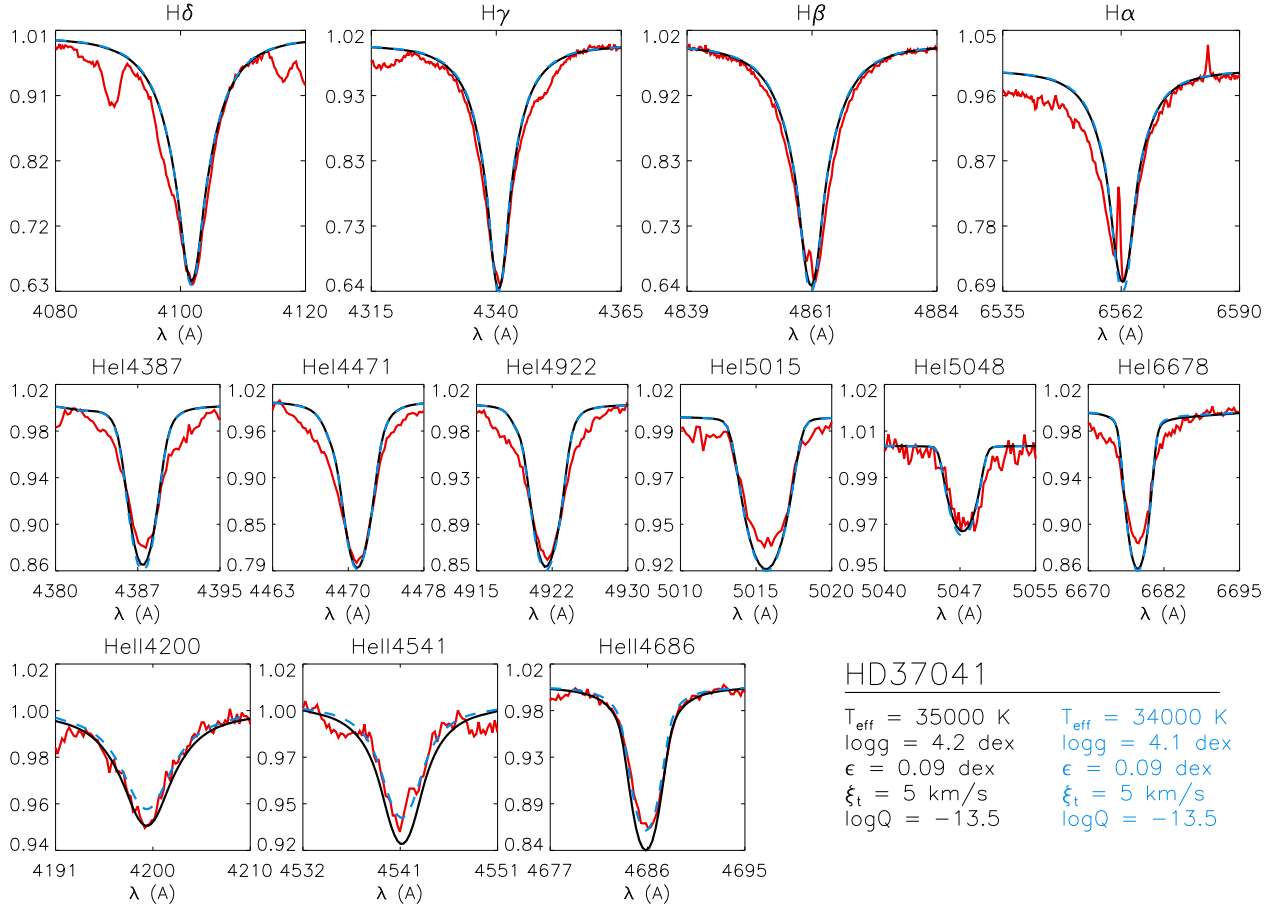


Fig. 6. As Fig. 4 for HD 37041 (θ^2 Ori A, O9V). A variation of 1000 K in T_{eff} along with 0.1 dex in $\log g$ has been considered in this case. Note that the broad wings of the He I cannot be fitted correctly (see text)

Star	T_{eff} (K)	$\log g$ (dex)
HD 37020	29970 / 30000	3.92 / 4.0
HD 37023	32600 / 32000	4.70 / 4.2
HD 37042	31600 / 29000	4.70 / 4.2

Table 5. Comparison of stellar parameters for HD 37020, HD 37023, HD 37042. First values refer to Cunha & Lambert (1992) determinations, second values to this work. We see that there is excellent agreement for HD 37020, but poor agreement (specially for $\log g$) for the other two stars.

Figure 5 shows the fitting of the HHe lines. Observed He I lines are slightly broader than the theoretical ones. Table 5 shows the comparison between the stellar parameters we have derived and those by Cunha & Lambert (1992). In this case the agreement is not so good as for HD 37020, although the T_{eff} are compatible, the $\log g$ they derived is very high.

HD 37041: This is a single-lined spectroscopic binary (see Howarth et al. 1997). These authors find a single peak in the cross correlation function of the IUE spectrum, indicating that the spectrum of the primary is not seriously contaminated.

Comparing the spectrum of this O9V star with that of the standard star 10 Lac (also classified as O9V) we have found that there are no unblended metal lines due to its high rotational velocity (except Si IV $\lambda 4089$, but it is in the blue wing of H δ). A good $v \sin i$ determination has been possible using the Fourier method with the He I lines. A $v \sin i = 131 \pm 4$ km s $^{-1}$ has been derived. We have used the Si IV line for checking the reliability of this value; a $v \sin i = 136 \pm 5$ km s $^{-1}$ is obtained. He II does not give good results. A $v \sin i = 140 \pm 11$ km s $^{-1}$ is derived using the FWHM method for the He I $\lambda 5015$ line.

Figure 6 shows the fitting of the synthetic profiles with the observed ones. A $v \sin i = 131$ km s $^{-1}$ has been considered for the H-He analysis. The wings of the He I lines cannot be well fitted; this might be explained by the presence of a companion.

HD 37042: We have found good agreement between metal and He I lines when using the Fourier method. The $v \sin i$ derived is 34.0 ± 0.5 km s $^{-1}$. The value we obtain with the FWHM method is 36 ± 3 km s $^{-1}$, in agreement with the former one.

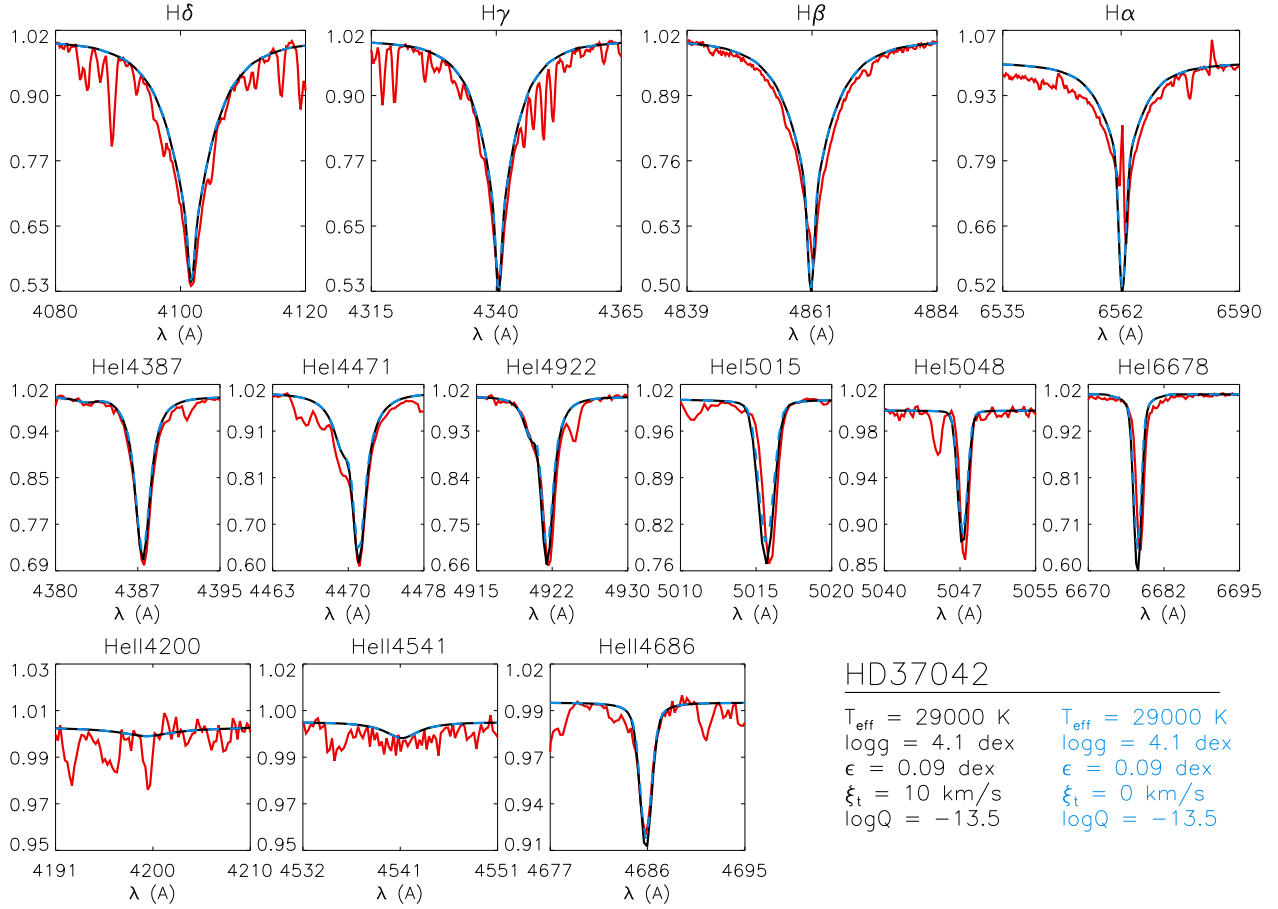


Fig. 7. As Fig. 4 for HD 37042 (θ^2 Ori B, B0.5V). For this star, the He I lines fit better if a microturbulence of 10 km s^{-1} is considered

A very good fit of the observed and synthetic FASTWIND profiles is obtained (see Figure 7). Again the forbidden component of He I $\lambda 4471$ is too weak. For this star, the He I lines fit better if a microturbulence of 10 km s^{-1} is considered.

This is the third star in common with Cunha & Lambert (see Table 5); we also find for this star (as for HD 37023) that the stellar parameters derived by these authors are very different from ours (they obtain a very high $\log g$ and a $T_{\text{eff}} \sim 3000 \text{ K}$ higher).

4.2. Comments on the reference stars

HD 214680: This star has been previously considered as standard for stellar parameter determination through the H-He analysis (see Herrero et al. 2002 and references therein). We have reanalysed here the spectrum of the star, obtained from a new observing run covering a larger spectral range (the He I $\lambda\lambda 5015$, 5048, 6678 and He II $\lambda 6683$ lines can be used for the new analysis). We are using a new FASTWIND version (Puls et al. 2005), slightly different from that used by Herrero et al. The new analysis will be useful for

checking the consistency of the latest version of the code in this stellar parameters regime.

The projected rotational velocity of this star has been determined accurately by means of the FWHM method ($v \sin i = 37 \pm 4 \text{ km s}^{-1}$). The Fourier method applied to the INT+IDS spectrum gives a $v \sin i = 30 \pm 0.8 \text{ km s}^{-1}$. This larger difference could be due to the fact that the $v \sin i$ is close to the computational Fourier transform limit ($\sim 20 \text{ km s}^{-1}$ for this spectrum), or because the microturbulence is affecting the determination of the rotational first zero in the Fourier transform (Gray 1973).

The whole set of HHe lines are perfectly fitted with the FASTWIND synthetic profiles (Figure 8). The parameters derived by Herrero et al. (2002) are $T_{\text{eff}} = 35500 \text{ K}$, $\log g = 3.95$ and $\epsilon = 0.09$. Our results are in agreement with those values.

HD 149438: This star was selected for comparison in the stellar oxygen abundance analysis of the B0.5V Trapezium cluster stars (see section 6). It has been studied elsewhere (Martin 2004; Przybilla & Butler 2004; Kilian 1994; Schönberner et al. 1988; Becker

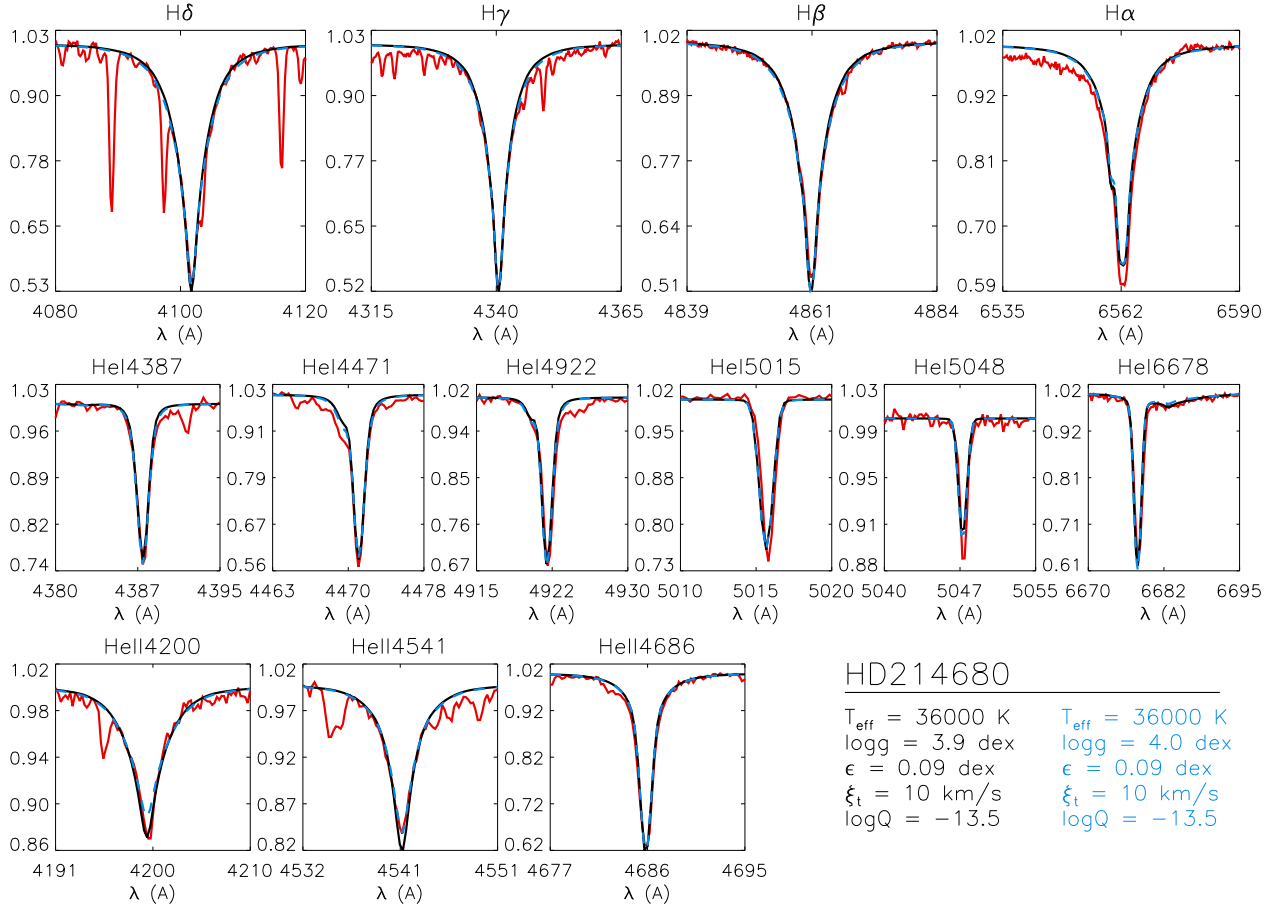


Fig. 8. As Fig. 4 for HD 214680 (10 Lac, O9V). A variation of 0.1 dex in $\log g$ has been considered in this case.

& Butler 1988; Peters & Polidan 1985; Kane et al. 1980). Comparing isolated oxygen and silicon lines to a set of rotationally broadened profiles, Kilian et al. (1991) obtained a $v \sin i = 19 \text{ km s}^{-1}$ for this star. The Fourier method applied to the CASPEC spectrum only allow us to say that the $v \sin i$ is lower than 13 km s^{-1} ; this is because in this case the effect of the microturbulence over the broadening of the metal lines can be comparable with that produced by the rotation, so the first zero could be associated with the microturbulence instead of the $v \sin i$ (Gray 1973).

A very good fit with the synthetic FASTWIND profiles is obtained for this star (see Figure 9). In this case the problem with the forbidden component of the He I $\lambda 4471$ can be clearly seen. Table 11 (Sect. 6.4) summarized the stellar parameters obtained in this and previous work.

HD 47839: This star was selected for comparison with the main ionizing source in Orion (θ^1 Ori C). Some comments on the analysis of this star and the comparison with θ^1 Ori C are presented in section 5. It was classified as O7V(f) by Walborn (1972). Gies (1993) pointed for the first time that this star is a spectroscopic binary with a period of 25 years. He

estimated a mass for the primary of $M = 34 M_{\odot}$, and suggested that the secondary has a spectral type O9.5Vn (with very broad lines).

We have used the spectrum of HD 214680 (O9V) convolved with a high $v \sin i$ ($\sim 350 \text{ km s}^{-1}$), for recognizing lines in the spectrum of HD 47839 not contaminated by the secondary star; three metal lines were found. Using these lines (Si IV $\lambda\lambda$ 4212, 4654 and N III λ 4379) a $v \sin i = 67 \pm 4 \text{ km s}^{-1}$ has been determined. A similar value has been derived applying the FWHM method to the same lines ($66 \pm 6 \text{ km s}^{-1}$).

The fitting of the H and He lines for the stellar parameters show how the He I are contaminated by the secondary star lines. The spectroscopic derived mass ($36 \pm 9 M_{\odot}$) is in very good agreement with the dynamical mass derived by Gies (1993). Herrero et al. (1992) derived a $T_{\text{eff}} = 39500 \text{ K}$, $\log g = 3.70$ and $\epsilon = 0.07$ for this star. Although we would expect to obtain a lower T_{eff} due to the inclusion of *line-blanketing* effects, the value we have obtained is slightly higher because there is also a large difference between the $\log g$ we derived (4.0) and the one by Herrero et al. (3.7). There is also another difference; we do not need a lower He abundance for fitting the He lines. This could be due to a binarity effect;

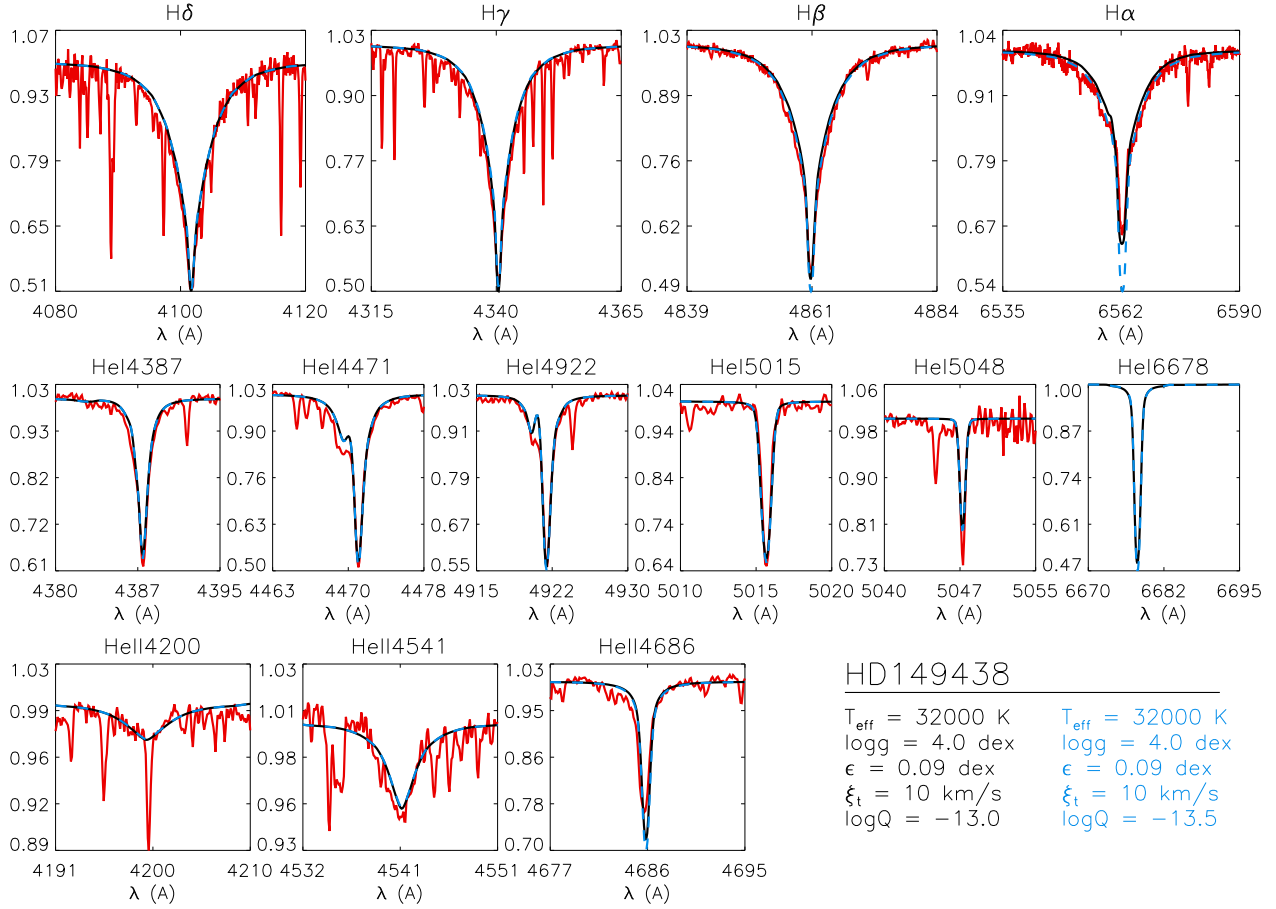


Fig. 9. As Fig. 4 for HD 149438 (τ Sco, B0.2V). A variation of 0.5 dex in $\log Q$ has been considered in this case. The H_α line is not contaminated by nebular emission for this star and a more accurate value of $\log Q$ can be determined. The He I λ 6678 is out of the observed range in the CASPEC spectrum of τ Sco.

when a composite spectrum is considered in a binary system, the lines can appear diluted or magnified due to the combination of the fluxes of the primary and the companion. If the system is out of eclipse the total flux will be higher than when the primary is observed isolated, so when the spectrum is normalized all the lines will appear diluted (and then a lower He abundance is needed to fit the He lines and a lower gravity is derived).

4.3. Results of the stellar parameters study

From the optical spectra of the Orion stars only upper limits for the Q parameter can be achieved. These estimations are based on the effect of the wind on the He II λ 4686 and H_α lines (the later one is contaminated by the nebular contribution). Some tests have shown that the other H and He lines are not affected by the uncertainties in $\log Q$ for the range of values considered, so the derived parameters will not be affected.

Masses, radii and luminosities have been derived for all the studied targets (they are indicated in Table 4 together with their uncertainties). The main source of un-

certainty for these parameters is that associated with the absolute magnitude (except for very large uncertainties in $\log g$). An error in $M_v \sim 0.3$ propagates to the mass, radius and logarithmic luminosity, giving uncertainties of $\sim 37\%$, 13% and 3% respectively.

The stars have been plotted on the HR diagram in Figure 11. Evolutionary tracks from Meynet & Maeder (2003), corresponding to initial masses ranging from 9 to $120 M_\odot$ and initial rotational velocities of 0 km s^{-1} are also plotted. All stars are found in the Main Sequence close to the ZAMS, as it is expected because of their youth. Nevertheless, we can see the separation from the ZAMS increasing with luminosity, as pointed out by Herrero et al. (2004). The loci of the Orion stars is consistent with an isochrone at about $2.5 \pm 0.5 \text{ Myr}$, derived from the tracks with zero initial rotational velocity, which is slightly older than the upper limit given by Palla & Stahler (1999, 2 Myr) and somewhat larger than other Trapezium age determinations found in the literature (e.g. Hillenbrand, 1997, $\leq 1 \text{ Myr}$). However, it has to be considered that, at large initial rotational velocities, the ZAMS is slightly shifted to the right and modifies the derived ages. Hence, until the role of the

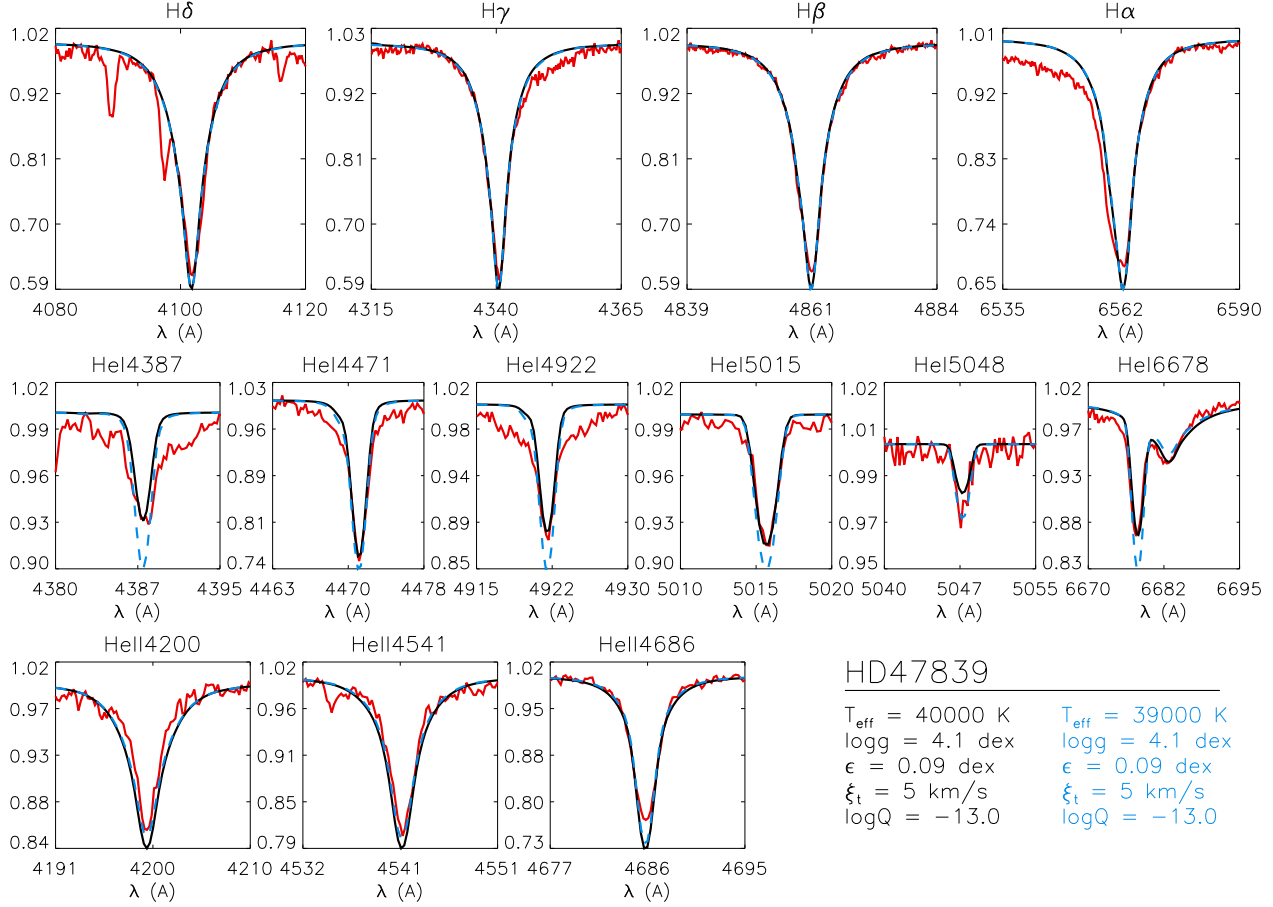


Fig. 10. As Fig. 4 for HD 47839 (15 Mon, O7V). A variation of 1000 K in T_{eff} has been considered in this case.

initial rotational velocities is properly understood (for example, the distribution of initial rotational velocities in clusters), the use of isochrones for massive stars in very young clusters should be regarded with special caution.

A good agreement between gravities derived from the evolutionary tracks and those obtained from quantitative analysis of the spectra is found (see Table 6). There is a trend for the most massive stars to have larger spectroscopic than evolutionary masses, but the number of objects is too small to draw any general conclusion.

5. Modeling the main ionizing star of the Orion nebula

HD 37022 (θ^1 Ori C, O7V) is the main ionizing source of the Orion nebula. We want to derive its stellar parameters as a first step to determine the effect of its ionizing flux on the photoionization of the surrounding nebula in a consistent way. Once these parameters are known, the spectral energy distribution can be modeled by using model atmosphere codes. In this way one of the inputs used in the photoionization codes will be fixed consistently.

Star	$M_{\text{spec}} (M_{\odot})$	$M_{\text{evol}} (M_{\odot})$	$\log g_{\text{spec}}$	$\log g_{\text{evol}}$
θ^1 Ori A	14	15	4.0	4.03
θ^1 Ori C	45	33	4.1	3.98
θ^1 Ori D	18	16	4.2	4.16
θ^2 Ori A	39	22	4.1	3.97
θ^2 Ori B	9	13	4.1	4.26
10 Lac	20	25	3.9	4.02
15 Mon	40	34	4.1	4.05
τ Sco	11	16	4.0	4.16

Table 6. Comparison of masses and gravities derived from the evolutionary tracks and from the quantitative analysis of the spectra. The quoted $\log g_{\text{evol}}$ values are given corrected to two decimal places to be consistent with the corresponding evolutionary masses. Note however that these are not an indication of the precision of these values, which we consider to be 0.1 dex.

5.1. A historical review

This star is known to have variable spectral features varying in phase or antiphase with a period of $15.422 \pm 0.002 \text{ d}$ (Stahl et al. 1993; Walborn & Nichols 1994; Stahl et al. 1996). These variable features were discovered after Conti (1972) showed for the first time that θ^1 Ori C has a variable inverted P-Cygni profile in the HeII 4686 line. Among them are H_{α} emission, variability in

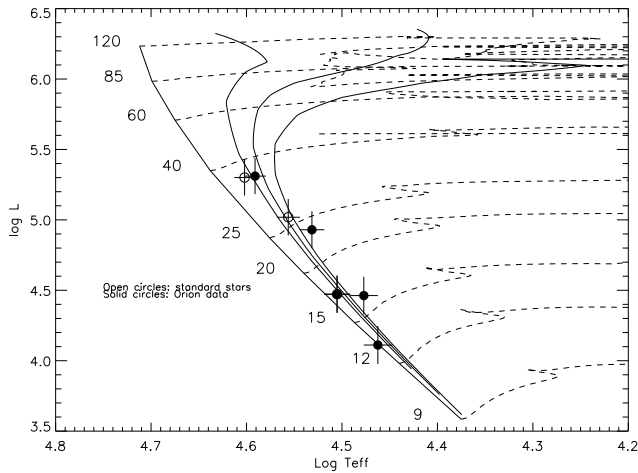


Fig. 11. HR diagram with the studied stars. Evolutionary tracks from Meynet & Maeder (2003). Isochrones from Schaller et al. (1992), corresponding to 2, 2.5 and 3 Myr

the equivalent width of some atmospheric and wind lines, and X-ray emission (Caillault et al. 1994; Gagné et al. 1997). Different explanations for this variability were postulated. The possibility of θ^1 Ori C being a binary and this binarity explaining the spectral variability has been dismissed (Stahl et al. 1996). The variability has been associated with the rotation of the star. Stahl et al. (1996) proposed the presence of a dipolar magnetic field in θ^1 Ori C, with the magnetic pole inclined 45° from the rotation axis (which is inclined 45° from the line of sight). The geometry of this system would imply alignment between magnetic pole and the line of sight at phase 0.5 and they would be perpendicular at phase 0.0 (when maximum H_α emission is found). Babel & Montmerle (1997) proposed the magnetically confined wind-shock model (MCWS) for explaining the variability in the star. According to this model, the radiatively line driven wind is confined by a dipolar magnetic field towards the magnetic equator of the system, generating a cold, dense disk due to the collision of material coming from both hemispheres, and a hot post-shock region.

The wind characteristics of θ^1 Ori C were determined by Howard & Prinja (1989) and Stahl et al. (1996) through UV spectrum studies. The former derived a mass loss rate of $4 \times 10^{-7} M_\odot \text{yr}^{-1}$, the latter through the absorption in C IV lines, determined a terminal velocity somewhat greater than 2500 km s^{-1} . It is common that O7V stars have stellar winds; what is not so common is the detection of the presence of magnetic fields in O stars. Donati et al. (2002) succeeded in the detection of Zeeman features in the spectrum of θ^1 Ori C through spectropolarimetric observations with the Anglo-Australian Telescope. They detected variability in the Stokes V profiles of some photospheric metal lines. This variability has a coherent modulation with the period derived from other variable features. However, the geometry derived

was in contradiction with that from Stahl et al. (1996), with the magnetic pole aligned with the line of sight at phase 0 (they found a maximum in the longitudinal component of the magnetic field at this phase).

5.2. Preliminary study of the spectrum

Having spectral variability, it is very important to understand the cause of this variability and to determine which lines are reliable for the stellar atmosphere modeling before comparing synthetic and observed H - He profiles. Preliminary work with the INT+IDS spectrum showed that a better spectral resolution was needed to apply the Fourier method for obtaining the $v \sin i$. This spectrum did not allow us to have neither a good enough sampling ($\Delta\lambda$), nor to carry on a study of the variability, so we decided to use some FEROS spectra with better quality and covering all variability phases (see Table 7), that are available in the ESO archive.

Through the study of the FEROS spectra observed in the different phases we have found all the variable spectral features described in Stahl et al. (1996) and some new ones:

- **He II 4686:** Broad emission appears in the blue wing of the line (the so-called inverted P-Cygni profile, with maximum at $\phi \sim 0$). Broad emission is also present in the red wing (maximum at $\phi \sim 0.5$, minimum at $\phi \sim 0.8$).
- **Balmer lines:** These lines are affected by the same broad emission features than those in He II $\lambda 4686$. The emission features are stronger in H_α and H_β than in the other Balmer lines
- **Metal and He I-II lines:** All the line strength varies in phase, being larger at $\phi \sim 0$.

Name	Date	$MJD-2,400,000.5$	ϕ
f07341+51	16/10/98	51102.31	0.180
f85221	26/07/99	51385.43	0.539
f03551+61	08/10/98	51094.28	0.659
f04711+21	10/10/98	51096.39	0.796
f15241	28/11/98	51145.37	0.972

Table 7. FEROS spectra used for the study of the spectral variability of θ^1 Ori C. All spectra have been downloaded from the ESO-FEROS database except f85221, kindly provided by O. Stahl. The different phases have been calculated from $MJD_0 - 2,400,000.5 = 48832.5$ (Stahl et al. 1996) and $P = 15.422$ days

This variability can be easily explained considering the model proposed by Stahl et al. (1996) and developed by Babel & Montmerle (1997). According to this model we would have an O7V star with a disc. The disc is produced by the confinement of the wind by a dipolar magnetic field through the magnetic equator. We will have a cool

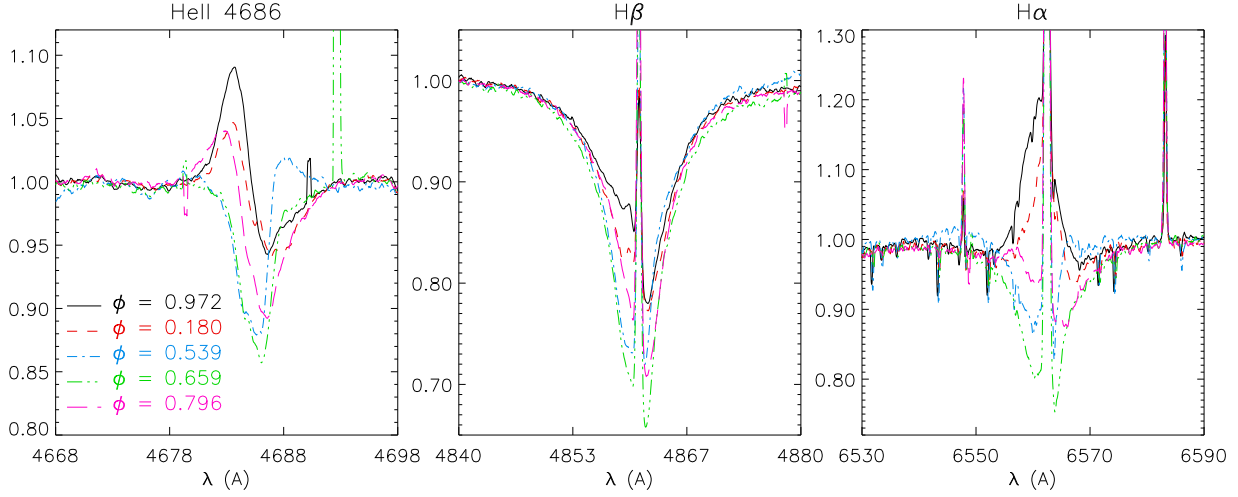


Fig. 12. The most representative phases have been selected for showing the variability of the He II $\lambda 4686$, H β and H α lines in the θ^1 Ori C spectrum. The variable feature associated with the inverted P-Cygni profile discovered by Conti in the He II $\lambda 4686$ line is also present in the hydrogen Balmer lines. Another emission feature can be clearly seen in the red wing of these lines at phase ~ 0.5 . The narrow emission features in the Balmer lines are nebular lines.

disc with material falling back to the stellar surface. If we consider that at phase 0 the cool disc is seen edge-on, the blue shifted emission appearing in He II $\lambda 4686$, H α and the other hydrogen Balmer lines can be explained as stellar photons absorbed and reemitted with a doppler shift corresponding to the velocity of the disc material falling to the surface of the star (in a process similar to what occurs in a stellar wind but with blue-shifted emission and red-shifted absorption). As density in the disc is very high then a strong blue-shifted emission will appear. At phase 0.5, when the disc is seen face-on the blue-shifted emission disappears. The emission appearing in the red wings of the former lines could be explained as the effect of the scattering of stellar photons by the wind material confined by the magnetic field and accreting to the disc (see figure 12). The disc will also have continuum emission that will affect the total continuum flux received from the star. This effect will be maximum when the disc is face-on because the emitting region is larger at this phase. The variability observed in He I, He II and metal lines (except for the emission in He I $\lambda 4686$) is only a consequence of this effect. As the total continuum flux is varying with the phase, the normalized spectrum will be affected. All absorption lines will be artificially weaker when the disc continuum emission is at maximum. This variability in the lines can allow us to estimate the amount of continuum flux that comes from the disc, and then the visual magnitude variability. By assuming that at phase 0.0 the lines are not affected by the continuum emission from the disc, we have scaled the spectra at the other phases to fit in the former spectrum; the scaling factor will be related to the ratio of visual fluxes (i.e. the difference in magnitudes) between the stellar component and the stellar+disc contribution. This can be seen in Figure 13. From this study we would expect a variability of ~ 0.16 magnitudes. Kukarkin et al. (1981), in their catalogue of suspected variable stars,

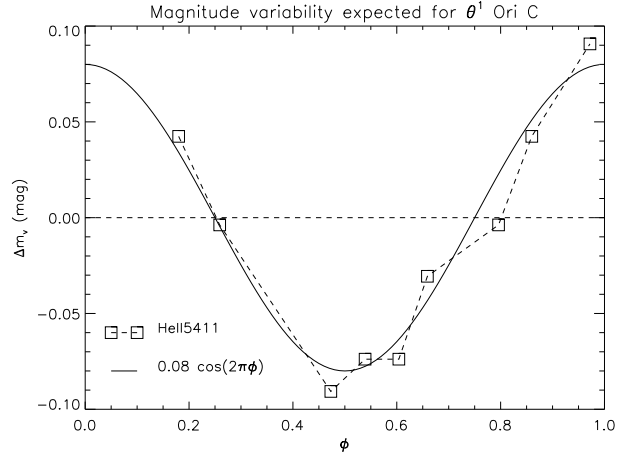


Fig. 13. Magnitude variability expected from the study of the He II $\lambda 5411$ line. The other He I-II lines follow a similar behavior. The presence of a disk could be responsible of this variability (see text for details). The solid line corresponds to a sinusoidal curve with a maximum change in m_v of 0.16, presented for comparison.

found a change in m_v of 0.06 magnitudes (5.10 - 5.16) for θ^1 Ori C. Hipparcos has also classified this star as variable; although Hipparcos data do not show a clear pattern, the median magnitude in Hipparcos system is 4.61 mag and the variability of this magnitude varies between 4.56 and 4.70, that is agreement with our study.

5.3. Determination of the $v \sin i$ of θ^1 Ori C

It is very important to have a good determination of the projected rotational velocity of this star as it is supposed that the spectral variability of θ^1 Ori C is related with its rotation. The derived $v \sin i$ should be independent on the phase and should be coherent with a period of

~ 15.4 days. It is shown in section 5.2 that the profiles of metal as well as He I and He II lines are dependent on the phase, however this is only an artificial dependence due to the presence of the disk. Once the spectra of different phases are corrected from the effect of dilution by the disc continuum emission, all the metal, He I and He II lines are independent on the phase (except those related with the disc, see Figure 12).

The Fourier method allows us to separate pure rotational broadening from other broadening mechanisms affecting the shape of the lines. We have used this method with some metal lines at phase $\phi \sim 0$ (see table 8). A $v \sin i = 24 \pm 3 \text{ km s}^{-1}$ has been derived. Figure 14 shows an example of the application of the Fourier method in the determination of the $v \sin i$ of θ^1 Ori C.

Table 7 offers a comparison of $v \sin i$ values obtained with the Fourier and FWHM methods. We see that the Fourier method gives more consistent values for all considered lines. Differences within the FWHM method may reach a factor of 2 and, in fact, we had problems when trying to fit the profile of some of the lines with a gaussian profile for measuring their linewidth.

The derived value for $v \sin i$ through the Fourier method is also more coherent with a O7V star rotating with a 15.422 days period. Considering $R \sim 11 R_{\odot}$, the upper limit for $v \sin i$ is $\sim 35 \text{ km s}^{-1}$ and therefore the inclination of the rotational pole is $i \sim 45$, in agreement with previous results (see next section).

Line	$v \sin i \text{ (km s}^{-1}\text{)}$	
	Fourier	FWHM
C III 4056	23	42 ± 10
N III 3998	23	28 ± 7
N IV 4057	20	33 ± 7
N IV 5200	20	32 ± 10
Si IV 4089	23	62 ± 14
Si IV 4654	30	47 ± 10
O III 4081	21	52 ± 10
O III 4376	26	31 ± 4
O III 4435	25	32 ± 5
O III 4454	27	34 ± 4
O III 5592	25	45 ± 6

Table 8. Projected rotational velocities derived through Fourier and FWHM methods for some metal lines present in the spectrum of θ^1 Ori C. Values were derived at phase 0.972 (see explanation in text).

5.4. Modeling of θ^1 Ori C

Once the spectral variability of θ^1 Ori C is understood and its $v \sin i$ has been derived, we can proceed to model its stellar atmosphere and wind through the observed spectrum of the star. The lines used for this analysis are shown in Figure 15; basically they are the ones used in

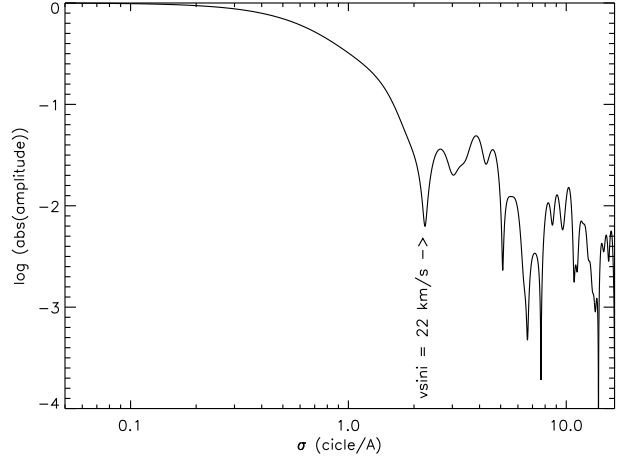


Fig. 14. Fourier analysis of the N IV $\lambda 4057$ line in θ^1 Ori C.

the other analyses plus He I 5875 Å. Some of the lines are contaminated (see Sect. 5.2), so this must be taken into account. The H_{δ} and H_{γ} lines have been selected as the most reliable lines for deriving $\log g$ (they are less contaminated than H_{β} and H_{α}); The whole set of He I-II lines has been considered except He II $\lambda 4686$; however it must be taken into account that the strength of all these lines vary with the phase. Phase 0.972 will be used for determining the stellar parameters (the effect of the continuum emission of the disc is smaller at this phase, see Sect. 5.2).

Although our study has shown that the rotational velocity (derived from Fourier analysis) is $\sim 24 \text{ km s}^{-1}$, when this broadening is considered all synthetic lines appear narrower than the observed ones. We have tried to solve the problem by means of a different $v \sin i$ value, however it does not work because then the shape of the synthetic profiles do not fit with the observed ones (the cores of the modeled profiles are too wide when the FWHM of the He lines is fitted). An extra-broadening mechanism has to be included. When a gaussian macroturbulent broadening (Gray, 1973) is used, the fit clearly improves, however in this case the H I and He II lines cannot be fitted simultaneously with the He I ones; when the former are fitted (for a $T_{\text{eff}} = 39000 \text{ K}$), some of the synthetic lines in the latter appear stronger and narrower than observed. A better fitting for the He I lines is obtained if a higher T_{eff} is considered, but then the synthetic He II lines appear too strong. There is no way to solve this problem in this region of the parameter space either by changing the rotational velocity, the macroturbulence or the microturbulence.

In section 4.2, figure 10 shows that the fitting of the He I-II lines could follow a similar behavior in the case of the O7V star HD 47839 (selected as reference star). A variation of 1000 K in the effective temperature change strongly the strength of the He I lines. For this spectrum

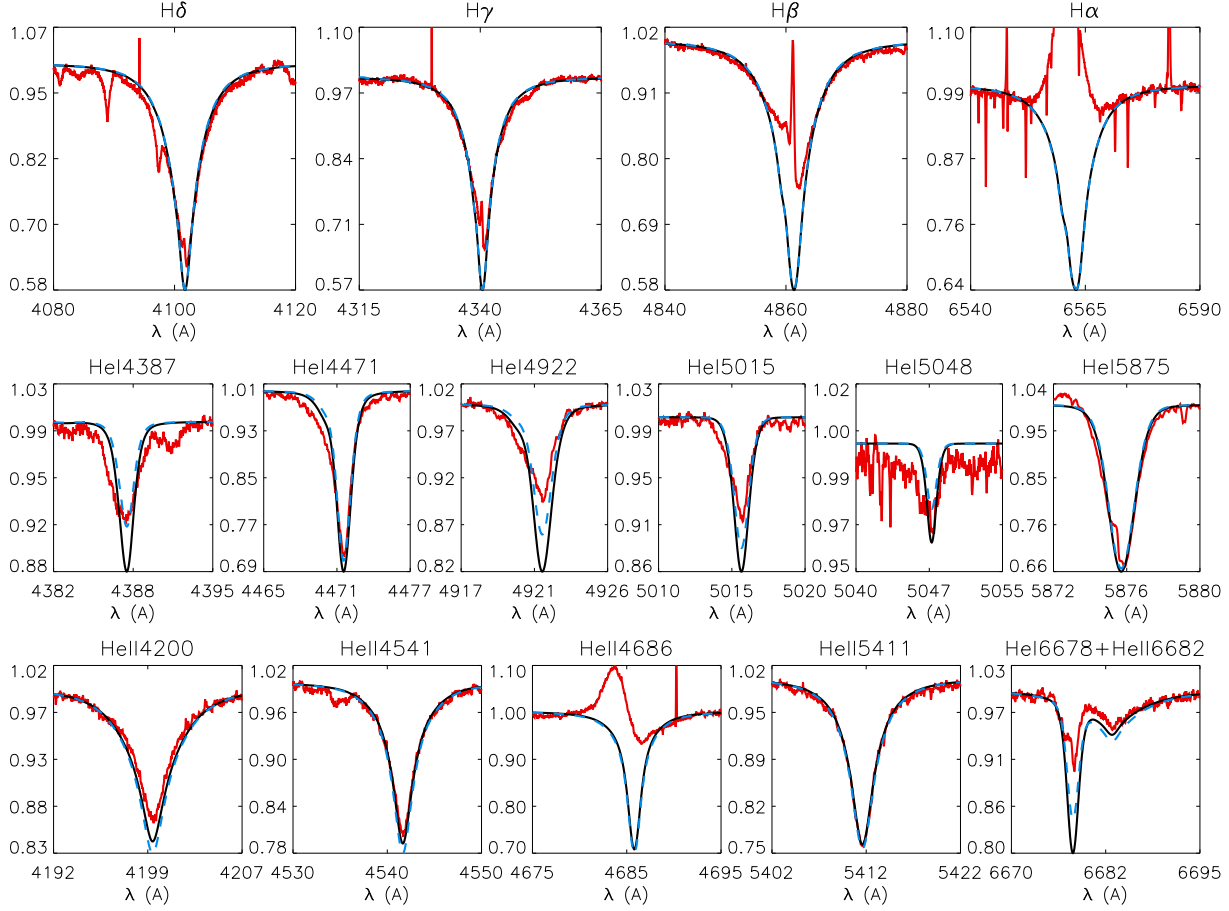


Fig. 15. HHe analysis of HD 37022 (θ^1 Ori C, O7Vp) at phase $\phi = 0.972$. The solid black line corresponds to a model with $T_{\text{eff}} = 39000$ K, $\log g = 4.1$, $\epsilon = 0.09$, $\xi_t = 5 \text{ km s}^{-1}$ and $\log Q = -13.0$; the dashed blue line corresponds to a model with $T_{\text{eff}} = 40000$ K (same remaining parameters). A $v \sin i = 24 \text{ km s}^{-1}$ and a macroturbulent velocity of 60 km s^{-1} have been considered. The observed lines strength has been enlarged by a factor 1.1 to correct for the effect of the disc continuum emission (see text for explanation).

also occurs that when the He I lines appear fitted, the He II lines are slightly stronger than observed, and if the He II lines are fitted, the He I lines are stronger than observed.

Puls et al. (2005) have shown that there is a discrepancy for the He I singlets between the synthetic FASTWIND and CMFGEN lines for effective temperatures between 36000 and 41000 K (being the CMFGEN profiles shallower). Therefore the He I triplet system can be considered more reliable (i.e. the He I $\lambda 4471$ line).

Knowing this discrepancy, we have considered the He I $\lambda 4471$ for the fitting with FASTWIND synthetic profiles. Our best model corresponds to $T_{\text{eff}} = 39000 \pm 1000$ K and $\log g = 4.1$ dex. From these values, the model flux and with $M_v = -4.9 \pm 0.3$, a stellar radius $R = 10.6 \pm 1.5 R_\odot$ is derived, and then an inclination of the rotational axis of $i = 44 \pm 12^\circ$. This value is in agreement with previous independent studies (Stahl et al. 1996, Donati et al. 2002), although our derived $v \sin i$ is more reliable and our radius is not obtained from a SpT - R calibration but is a result of the spectral analysis of the star.

6. Oxygen abundances in the Trapezium cluster B0.5V stars

Three of the stars studied in Orion are perfect targets for a stellar abundance analysis as they have many narrow unblended lines. These are HD 37020, HD 37023 and HD 37042, three B0.5V stars. Cunha & Lambert (1992, 1994) presented carbon, nitrogen, oxygen and silicon abundances from LTE and NLTE analyses for these stars. For comparison purposes, a similar analysis has been done for τ Sco, a B0.2V star (Walborn & Fitzpatrick, 1990) with very low $v \sin i$. Stellar abundances for this star have been derived elsewhere in the literature (Hardorp & Scholz 1971, Kane et al. 1980, Peters & Polidan 1985, Schönberner et al. 1988, Becker & Butler 1988, Kilian et al. 1991, Martin 2004, see Table 11). The other two Orion stars have been ruled out: HD 37041 has a relatively high projected rotational velocity, so metallic lines are broadened and then appear blended; HD 37022, being an O7V star, does not have enough oxygen lines for an accurate abundance analysis.

We have therefore derived oxygen abundances for the Trapezium cluster B0.5V stars for comparing them with the M42 nebular abundances obtained by Esteban et al. (2004).

6.1. Line identification and measurement of equivalent widths

We make use of the classical method of curve of growth to determine the oxygen abundances. When this methodology is used it is important to remove all lines that appear blended. The spectrum of τ Sco (a star with similar spectral type than our targets in Orion and a very low $v \sin i$) has been used for an identification of the O II lines present in the spectra of the Orion stars. The whole set of lines is shown in Table 13 divided into multiplets, together with them $\log gf$ values (basically they are taken from the NIST database).

The equivalent widths of all the O II lines listed in Table 13 have been measured for the four stars, however only our set of suitable lines (see below) is shown in Table 9. To measure the equivalent widths we use our own software developed in IDL. A least squares profile fitting procedure was used, with Gaussian profiles fitting the line and polynomials of degree one or two to fit the local continuum. Errors in the measurements due to the uncertainty in the position of the local continuum (estimated as $\pm 1/SNR$, Villamariz et al. 2002) have also been considered. The estimated value of the uncertainty in the measurement of the EWs is ~ 5 mÅ and ~ 10 mÅ for some problematic lines.

6.2. Line selection

Some of the lines that appear unblended in the spectrum of τ Sco cannot be used in the analyses of the other stars; they have larger rotational broadening and then appear blended (or lie in the wings of H or He lines). Once a first set of unblended lines was selected, a preliminary abundance analysis was done separately for each multiplet. In this way the dispersion in the line abundances for the zero slope value of the microturbulence are minimized as all the lines in a multiplet are formed in the same region in the stellar atmosphere. Problematic lines, errors in the measurement of the equivalent widths or artificial trends can be then identified; such lines will be removed in the global analysis (e.g. this is the case for the O II $\lambda 4414$ line, an isolated line that gives systematically lower abundances).

The set of suitable lines finally used in the abundance analyses is presented in Table 9. They are lines coming from transitions between configurations $2p^2(^3P)3s-2p^2(^3P)3p$ (NIST multiplets 64, 65 and 72) and $2p^2(^3P)3p-2p^2(^3P)3d$ (NIST multiplets 90, 148 and 130). We have ruled out the lines that do not follow the general trend. We have found that lines from

multiplets 99, 118, 161, 188, 172 give systematically lower abundances. This effect can be due to the definition of the oxygen model atom we have used, or can be associated with the $\log gf$ values. Some preliminary comparisons with TLUSTY analyses have shown that the difference in the line abundances is also present for the case of τ Sco.

6.3. Chemical analysis

For each star we proceed as follows. A grid of 16 FASTWIND models combining four abundances and four values of microturbulence is calculated. In this way, the curves of growth for each line can be constructed by plotting the theoretical equivalent width for each abundance and microturbulence versus the abundance (see Figure 16 for the case of the O II $\lambda 4414$ line in HD 37042).

Through the observed equivalent width and its error one abundance (and its derived uncertainty) can be derived for each line and microturbulence. The individual line abundances are dependent on the microturbulence. The microturbulence value that minimizes the dependence of the line abundances on the line strength in the $\log \epsilon$ -EW diagrams will be the microturbulence we are looking for. These diagrams can be also used as a diagnostic tool to check the reliability of the different lines for the abundance determination (see Section 6.2). Figure 17 shows the $\log \epsilon$ -EW diagrams for two different microturbulences in the study of HD 37023.

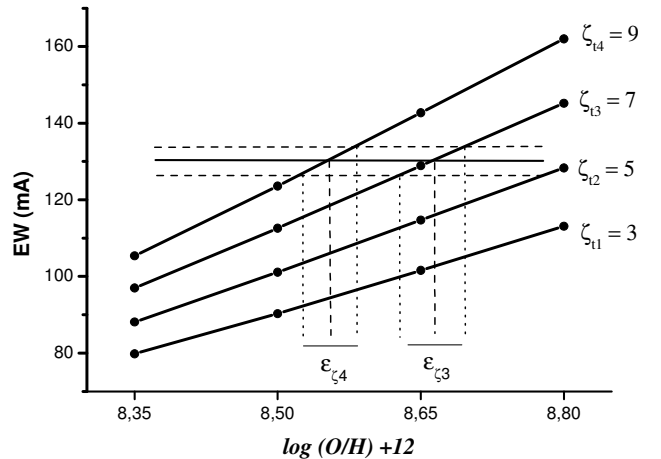


Fig. 16. Example of curve of grow for the line O II $\lambda 4414$ in the star HD 37042. A grid of 16 models has been considered (4 microturbulences and 4 oxygen abundances). The observed EW of the line and its uncertainty are plotted as horizontal lines. Two examples of abundances (and their uncertainties) derived for microturbulences 7 and 9 km s^{-1} are plotted as vertical lines.

The microturbulence derived from the zero slope for each star is presented in Table 10. Uncertainties in the microturbulence are obtained considering the errors derived

Line	EW_o	$\epsilon(O)^a$	$\Delta\epsilon(O)$	EW_o	$\epsilon(O)$	$\Delta\epsilon(O)$	EW_o	$\epsilon(O)$	$\Delta\epsilon(O)$	EW_o	$\epsilon(O)$	$\Delta\epsilon(O)$			
HD 149438 ($\xi_t = 8.7$)				HD 37020 ($\xi_t = 6.5$)			HD 37023 ($\xi_t = 7.4$)			HD 37042 ($\xi_t = 5.5$)					
$T_{\text{eff}} = 32000 \text{ K}$		$\log g = 4.0$		$T_{\text{eff}} = 30000 \text{ K}$		$\log g = 4.0$		$T_{\text{eff}} = 32000 \text{ K}$		$\log g = 4.2$		$T_{\text{eff}} = 29000 \text{ K}$		$\log g = 4.2$	
O II 4638	85	8.75	0.04	—	—	—	—	—	—	—	—	—	—	—	—
O II 4641	127	8.65	0.04	—	—	—	—	—	—	—	—	—	—	—	—
O II 4661	90	8.70	0.04	80	—	—	76	8.60	0.06	88	8.56	0.06	—	—	—
O II 4676	79	8.71	0.05	82	8.59	0.06	65	8.59	0.06	82	8.63	0.06	—	—	—
O II 4317	92	8.69	0.03	107	8.67	0.04	78	8.59	0.05	98	8.60	0.07	—	—	—
O II 4319	86	8.71	0.04	100	8.69	0.05	75	8.63	0.05	95	8.66	0.07	—	—	—
O II 4366	80	8.61	0.04	96	8.62	0.05	77	8.62	0.05	103	8.62	0.08	—	—	—
O II 4416	100	8.62	0.04	120	8.65	0.05	93	8.62	0.05	117	8.70	0.11	—	—	—
O II 4452	37	8.67	0.07	48	8.64	0.07	32	8.58	0.09	60	8.67	0.10	—	—	—
O II 4072	—	—	—	118	8.61	0.12	99	8.57	0.09	119	8.71	0.10	—	—	—
O II 4076	—	—	—	133	8.59	0.12	114	8.57	0.09	—	—	—	—	—	—
O II 4078	—	—	—	46	8.67	0.17	33	8.59	0.09	48	8.68	0.10	—	—	—
O II 4941	38	8.64	0.06	50	8.68	0.07	35	8.60	0.09	43	8.60	0.08	—	—	—
O II 4943	60	8.62	0.04	67	8.58	0.06	55	8.59	0.07	56	8.48	0.07	—	—	—
O II 4956	17	8.73	0.12	20	8.67	0.14	13	8.58	0.20	22	8.71	0.13	—	—	—
O II 4891	24	8.75	0.10	27	8.68	0.11	18	8.59	0.15	27	8.65	0.12	—	—	—
O II 4906	34	8.64	0.07	42	8.63	0.08	31	8.59	0.10	47	8.72	0.08	—	—	—
$\epsilon(O) = 8.70 \pm 0.10$				$\epsilon(O) = 8.65 \pm 0.10$			$\epsilon(O) = 8.59 \pm 0.10$			$\epsilon(O) = 8.64 \pm 0.10$					

$$^a \epsilon(O) = \log(O/H) + 12$$

Table 9. Equivalent widths and derived line abundances for the set of O II lines used in our analysis. Line abundances refer to the microturbulence given in brackets for each star (ξ_t in km s^{-1}). Uncertainties in the line abundances come from the propagation of the uncertainties of the equivalent width measurements (see text). Some of the O II lines of the Orion stars have not been used as they appear blended. O II $\lambda\lambda$ 4072, 4076 and 4078 lines were ruled out in the analysis of τ Sco due to the poor quality of the CASPEC spectrum in this region. Final oxygen abundances for each star have been calculated through a weighted mean of the linear values. Errors represent the statistical deviation for these mean values.

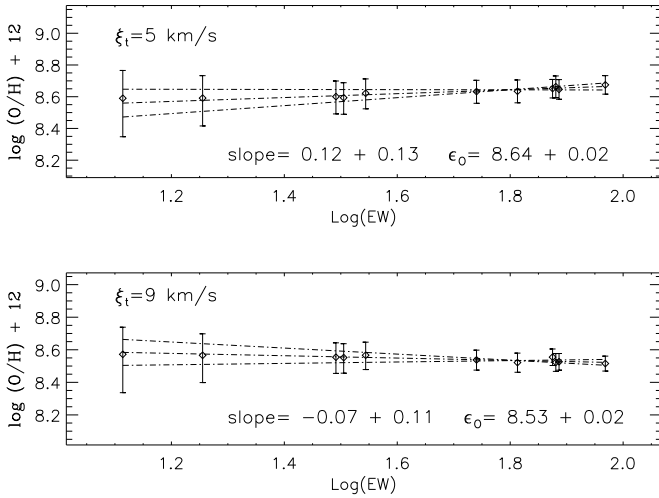


Fig. 17. Example of $\log \epsilon$ - EW diagrams in the study of HD 37023.

for the slope in a linear fit of the data (due to errors in the individual line abundances). This step also allows us to estimate the contribution of the uncertainty in the microturbulence to the the total oxygen abundance. This uncertainty depends mainly on the quality of the spectra (in this case is ~ 0.06 dex).

Abundance values for each line as well as their uncertainties are obtained for that microturbulence (see Table 9). The final abundance value is calculated through

a weighted mean of the linear individual line abundances ($10^{\epsilon_i - 12}$), and its uncertainty (σ_w) is that associated with this mean. Final values for the total oxygen abundances for each star are shown in Table 10. The final uncertainty in the oxygen abundance takes into account four different sources of errors: those associated with the statistical analysis, those derived from the error in the determined microturbulence and finally those referred to the uncertainty in the stellar parameters and to atomic data. All these sources of error are added quadratically for deriving the final abundance uncertainty (see Villamariz et al. 2002).

6.4. Oxygen abundance in the standard star τ Sco

Table 11 summarizes the oxygen abundances appeared in the literature for τ Sco. Our derived value is compatible with previous results but a little higher than most of them. This difference can be easily explained taking into account that for this range of stellar parameters, the oxygen abundance derived through O II lines is very sensitive to a change in T_{eff} and $\log g$ (the lines become fainter when a higher T_{eff} is considered and then the derived oxygen abundance is higher). The uncertainty in the oxygen abundance due to a change of $\sim 1000 \text{ K}$ in T_{eff} is ~ 0.08 dex. This effect is considered in the uncertainty for the given values, however the central value will be slightly dependent on the derived stellar parameters. We have considered two FASTWIND models with different T_{eff}

Object	N	ξ_t (km/s)	$\log(\text{O}/\text{H})+12$
This work			
HD 149438	14	8.1 ± 1.2	8.70 ± 0.10
HD 37020	13	6.8 ± 1.3	8.65 ± 0.10
HD 37023	15	6.3 ± 3.0	8.59 ± 0.10
HD 37042	13	4.9 ± 2.0	8.64 ± 0.10
Cunha & Lambert (LTE)			
HD 37020	6	7.0	8.92 ± 0.05
HD 37023	7	7.5	8.76 ± 0.04
HD 37042	7	6.0	8.97 ± 0.04
Cunha & Lambert (NLTE)			
HD 37020	7	5.0	8.83 ± 0.12
HD 37023	6	7.0	8.87 ± 0.08
HD 37042	6	6.0	8.85 ± 0.06
Esteban et al.			
Gas			8.65 ± 0.03
Gas + dust			8.73 ± 0.03

Table 10. Oxygen abundances for the three B0.5V stars inside Orion Nebula and the reference star τ Sco. Oxygen LTE and NLTE abundances derived by Cunha & Lambert (1994) for the Orion stars as well as those calculated by Esteban et al. (2004) for the nebula are also presented for comparison.

(32000, 32500 K) and the same $\log g$ obtaining oxygen abundances of 8.70 and 8.74 dex respectively.

The derived abundance is also very dependent on the microturbulence considered (specially if lines with high equivalent width are used). We have taken into account these dependencies in our uncertainties.

Work	T_{eff} (K)	$\log g$	$\epsilon(\text{O})$	ξ_t (km s $^{-1}$)
Hardorp & Scholz	30900	4.05	8.70	—
Kane et al.	30300	3.95	8.6	—
Peters & Polidan	31500	4.3	8.65 ± 0.26	5
Schönberner et al.	33000	4.15	8.5	5
Becker & Butler	33000	4.15	8.70 ± 0.19	5
			8.58 ± 0.19	10
Kilian et al.	31400	4.24	8.61 ± 0.12	3
Martin	30000	3.75	8.58 ± 0.17	7
This work	32000	4.0	8.70 ± 0.10	8

Table 11. Comparison of stellar parameters and abundances derived for τ Sco in previous studies found in the literature and in this work.

6.5. Comparing with nebular and previous stellar results

The oxygen abundances derived for the Orion stars are compatible within the errors (see Table 10). HD 37023 has a slightly lower abundance (but still compatible with the other abundances). In section 4.1 we have seen that the fitting of the H and He lines is not so good as for the other B0.5V stars (observed lines appear slightly broader).

The oxygen abundances in the Orion stars are systematically lower than that derived for τ Sco.

Esteban et al. (2004) have recently published a reappraisal of the chemical composition of the Orion nebula. They derived a total oxygen gas-phase abundance $\epsilon(\text{O}) = 8.65 \pm 0.03$. However, some oxygen is expected to be depleted onto dust grains in ionized nebula, so the total gas+dust oxygen abundance should take this depletion into account. In a previous work (Esteban et al. 1998), these authors estimate the depletion onto dust grains by comparing Si and Fe nebular abundances with those obtained by Cunha & Lambert (1994) for B stars in the Orion association, assuming a certain composition for the main dust molecules. Taking into account this correction, the final gas+dust oxygen abundance Esteban et al. (2004) propose is $\epsilon(\text{O}) = 8.73 \pm 0.03$ (an oxygen abundance correction for dust ~ 0.08 is applied).

Our stellar results are compatible with those obtained by Esteban et al. (2004) for the gas phase, however the dust+gas corrected abundance seems to be too high compared with our derived stellar abundances, although still marginally consistent inside the uncertainties.

Our oxygen abundances are systematically lower than the NLTE abundances by Cunha & Lambert (1994). In that paper the former authors comment that the LTE abundances are slightly more reliable than the NLTE abundances they present. The difference between their LTE abundances and our results are even higher. These differences can be in part associated with the differences in the derived stellar parameters for these stars. The T_{eff} and $\log g$ obtained by these authors are higher than ours (see Table 5), so the oxygen abundances they derive are obviously higher.

This difference in the effective temperatures may also affect the derived silicon abundances (used by Esteban et al. 1998 for estimating the oxygen depletion). Cunha & Lambert (1994), derive their silicon abundances by using 3 Si III lines. Preliminary silicon analysis by our group has shown that a difference of ~ 1000 K in T_{eff} can shift the Si III abundances up to 0.2 dex (deriving a higher abundance for the lower T_{eff}).

Alternatively, our result could suggest that the molecules that Esteban et al. used to estimate the O dust depletion in Orion cannot be present in large amounts in this nebula. Consequently, Si, Mg and Fe (the refractory elements being the main constituents of those molecules) have to form other molecules without oxygen.

7. Conclusions

By means of a detailed spectroscopic analysis of the optical spectra of the Trapezium cluster stars, stellar parameters and oxygen abundances have been derived.

Label	λ (Å)	$\log gf$	Notes
$2s^2 2p^2$ (3P) 3p - $2s^2 2p^2$ (3P) 3d			
NIST 90 (4D ₀ -4F)			
O II 4069	4069.623	0.149	+ C III
	4069.882	0.344	
O II 4072	4072.153	0.552	
O II 4076	4075.862	0.693	+ C II
O II 4078	4078.842	-0.284	
NIST 118 (2D ₀ -2F)			
O II 4699	4699.218	0.269	
	(4699.011	0.418)	NIST 172
	(4698.437	-0.883)	NIST 172
O II 4705	4705.346	0.477	
O II 4741	4741.704	-0.987	Very weak
NIST 130 (4S ₀ -4P)			
O II 4891	4890.856	-0.436	
O II 4906	4906.830	-0.160	
NIST 148 (2P ₀ -2D)			
O II 4941	4941.072	-0.054	
O II 4943	4943.005	0.239	
O II 4956	4955.707	-0.573	Very weak
$2s^2 2p^2$ (3P) 3s - $2s^2 2p^2$ (3P) 3p			
NIST 64 (4P-4D ₀)			
O II 4638	4638.856	-0.332	+ C II + Si III
O II 4641	4641.810	0.054	+ N III
O II 4650	4649.135	0.308	
O II 4651	4650.638	-0.362	
O II 4661	4661.632	-0.278	
O II 4673	4673.733	-1.088	
O II 4676	4676.235	-0.394	
O II 4696	4696.352	-1.380	
NIST 65 (4P-4P ₀)			
O II 4317	4317.139	-0.386	
	4317.696	...	
O II 4319	4319.630	-0.380	+ N III
	4319.866	-0.502	
O II 4366	4366.895	-0.348	
	4366.530	...	
NIST 72 (2P-2D ₀)			
O II 4414	4414.899	0.172	
	4414.456	-1.483	
O II 4416	4416.975	-0.076	+ Si IV
O II 4452	4452.378	-0.789	

Table 12. Preliminary set of O II lines selected for the analysis, divided by multiplets. The spectrum of the low $v \sin i$ star τ Sco, has been used to identify the lines. $\log gf$ values are from the NIST database.

Projected rotational velocities has been obtained through Fourier method. This method has been extensively used for late type stars, but not widely applied to early type stars. Our results show this method to be very useful for distinguishing between rotational broadening and another broadening mechanisms that can be present in early type stars (i.e. macroturbulence). The agreement is very good when comparing with results from the line-width method. The Fourier method applied to the high resolution θ^1 Ori C FEROS spectra, allow us to derive a very accurate $v \sin i$ that is in agreement with the period of variability of some spectral features in θ^1 Ori C.

Label	λ (Å)	$\log gf$	Notes
$2s^2 2p^2$ (1D) 3s - $2s^2 2p^2$ (1D) 3p			
NIST 99 (2D-2F ₀)			
O II 4590	4590.974	0.350	
O II 4596	4595.957	-1.032	
	4596.177	0.200	
$2s^2 2p^2$ (1D) 3p - $2s^2 2p^2$ (1D) 3d			
NIST 161 (2F ₀ -2G)			
O II 4185	4185.440	0.604	+ C III
O II 4189	4189.788	0.716	
	4189.581	-0.828	
NIST 188 (2P ₀ -2P)			
O II 4691	4690.888	-0.610	
	4691.419	-0.309	
O II 4701	4701.179	0.088	
	4701.712	-0.611	
	4700.441	-3.298	
NIST 172 (2D ₀ -2F)			
O II 4703	4703.161	0.263	

Table 13. (Cont) Preliminary set of O II lines selected for the analysis, divided by multiplets. The spectrum of the low $v \sin i$ star τ Sco, has been used to identify the lines. $\log gf$ values are from the NIST database.

Stellar parameters and their uncertainties have been derived for the studied stars using H, He I and He II lines and FASTWIND code.

The presence of many O II lines in the optical spectrum of three B0.5V Orion stars has allowed us to work on a very detailed abundance analysis using the curve of growth method. This analysis has been performed through a careful selection of suitable lines from a previous study of the different O II multiplets. In this way, the dispersion in the line abundances is reduced, and the final abundance value derived is very precise.

The derived oxygen abundances in the Orion stars are in agreement with the nebular gas-phase abundances obtained by Esteban et al. (2004), and ~ 0.2 dex lower than the NLTE abundances derived by Cunha & Lambert (1994).

The gas+dust corrected oxygen abundances estimated by Esteban et al. (1998, 2004) using the Cunha & Lambert stellar abundances in the Orion association, seem to be too high compared with our derived abundances (although still marginally consistent within the uncertainties). This result suggests a lower dust depletion factor of oxygen than previous estimations for the Orion nebula. A revision of the silicon, magnesium and iron stellar abundances in the Trapezium cluster stars is presently under way in our group for confirming this result.

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References

- Babel, J., & Montmerle, T. 1997, *ApJ*, 485, L29
 Becker, S.R., & Butler K. 1988, *A&A*, 201, 232
 Caillault, J.P., Gagne, M., & Stauffer J.R. 1994, *ApJ*, 432, 386
 Carroll, J.A. 1933, *MNRAS*, 93, 478
 Conti, P.S. 1972, *ApJ*, 174, L79
 Crowther, P.A., Hillier, D.J., Evans, C.J., et al. 2002, *ApJ*, 579, 774
 Cunha, K., & Lambert, D.L. 1992, *ApJ*, 399, 586
 Cunha, K., & Lambert, D.L. 1994, *ApJ*, 426, 170
 Donati, J.F., Babel, J., Harries, T.J., et al. 2002, *MNRAS*, 333, 55
 Esteban, C., Peimbert, M., Torres-Peimbert, S., et al. 1998, *MNRAS*, 295, 401
 Esteban, C., Peimbert, M., García-Rojas, J., et al. 2004, *MNRAS*, 355, 229
 Ferland, G.F. 2001, *PASP*, 113, 165
 Gagné, M., Caillault, J.P., Stauffer, J.R., et al. 1997, *ApJ*, 478, L87
 Gies, D.R., Mason, B.D., Hartkopf, W.I., et al. 1993, *AJ*, 106, 2072
 Gray, D.F. 1973, *ApJ*, 184, 461
 Hardorp, J., Scholz, M. 1970, *ApJS*, 19, 193
 Herrero, A., Kudritzki, R. P., Vilchez, J. M., et al. 1992, *A&A*, 261, 209
 Herrero, A., Puls, J., & Najarro, F. 2002, *A&A*, 396, 949
 Herrero, A., Simón-Díaz, S., Najarro, F., et al. 2004, proceedings of the International Workshop on Massive Stars in interacting binaries (2004, Quebec, Canada)
 Hillenbrand, L.A. 1997, *AJ*, 113, 1733
 Hillier, D.J., & Miller, D.L. 1998, *ApJ*, 496, 407
 Howarth, I.D., Siebert, K.W., Hussain, G.A.J., et al. 1997, *MNRAS*, 284, 265
 Howarth, I.D., & Prinja, R.K. 1989, *ApJS*, 69, 527
 Hubeny, I., & Lanz, T. 1995, *ApJ*, 439, 875
 Humphreys, R. 1978, *ApJSS*, 38, 309
 Kane, L., McKeith, C.D., & Dufton, P.L. 1980, *A&A*, 84, 115
 Kilian, J., Becker, S.R., Gehren, T., et al. 1991, *A&A*, 244, 419
 Kilian, J. 1994, *A&A*, 282, 867
 Kukarkin, B.V., Kholopov, P.N., Artiukhina, N.M., et al. 1981, *CSV*
 Kurucz, R.L. 1979, *ApJS*, 40, 1
 Lohsen, E. 1975, *IBVS*, 988, 1
 Maeder, A., & Meynet G. 2000, *ARA&A*, 38, 143
 Massey, P., Bresolin, F., Kudritzki, R.P., et al. 2004, *ApJ*, 608, 1001
 Massey, P., Puls, J., Pauldrach, A.W.A., et al. 2005, *ApJ*, accepted, astro-ph/0503464
 Martin, J.C. 2004, *AJ*, 128, 2474
 Martins, F., Schaerer, D., & Hillier D.J. 2002, *A&A*, 382, 999
 Martins, F., Schaerer, D., & Hillier D.J. 2005, *A&A*, accepted, astro-ph/0503346
 McNamara, D.H., & Larsson H.J. 1962, *ApJ*, 135, 748
 Meynet, G., & Maeder, A. 2003, *A&A*, 404, 975
 O'Dell, C.R. 2001, *PASP*, 113, 290
 Palla, F. & Stahler, S.W. 1999, *ApJ*, 552, 772
 Pauldrach, A.W.A., Hoffmann, T.L., & Lennon, M. 2001, *A&A*, 375, 161
 Peters, G.J., & Polidan, R.S. 1985, *IAUS*, 111, 417
 Preibisch, T., Balega, Y., Hofmann, K.H., et al. 1999, *NewA*, 4, 531
 Przybilla, N., & Butler, K. 2004, *ApJ*, 609, 1181
 Puls, J., Urbaneja, M.A., Venero, R., et al. 2005, *A&A*, in press
 Repolust, T., Puls, J., & Herrero, A. 2004, *A&A*, 415, 349
 Roberto, M., Beckwith, S.V.W., & Panagia, N. 2004, *AJ*, in press
 Royer, F., Gerbaldi, M., Faraggiana, et al. 2002, *A&A*, 381, 105
 Santolaya-Rey, A.E., Puls, J., & Herrero A. 1997, *A&A*, 323, 488
 Schonberger, D., Herrero, A., Becker, S., et al. 1988, *A&A*, 197, 209
 Simón-Díaz, S., Herrero, A., & Esteban, C. 2003, *RMxAC* 18, 123
 Simón-Díaz, S., & Herrero, A., in preparation
 Smith, M.A., & Gray, D.F. 1976, *PASP*, 88, 809
 Schaller, G., Schaerer, D., Meynet, G., & Maeder, A. 1992, *A&AS*, 96, 269
 Stahl, O., Wolf, B., Gäng, Th., et al. 1993, *A&A*, 274, L29
 Stahl, O., Kaufer, A., Rivinius, T., et al. 1996, *A&A*, 312, 539
 Trundle, C., Dufton, P. L., Lennon, D. J., et al. 2002, *A&A*, 395, 519
 Urbaneja, M. A., Herrero, A., Bresolin, F., et al. 2005, *ApJ*, in press
 Vacca, W.D., Garmany, C.D., & Shull, J.M. 1996, *ApJ*, 460, 914
 Villamariz, M.R., & Herrero, A. 2000, *A&A*, 357, 597
 Villamariz, M.R., Herrero, A., Becker S. R., et al. 2002, *A&A*, 388, 940
 Walborn, N.R., & Fitzpatrick, E.L. 1990, *PASP*, 102, 379
 Walborn, N.R., & Nichols, J.S. 1994, *ApJ*, 425, L29
 Walborn, N.R. 1972, *AJ*, 77, 312